Accuracy of Bias Measurements from 3-point Galaxy Correlations

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arXiv: 1607.01024, 1503.00313, 1504.02074, 1403.1259

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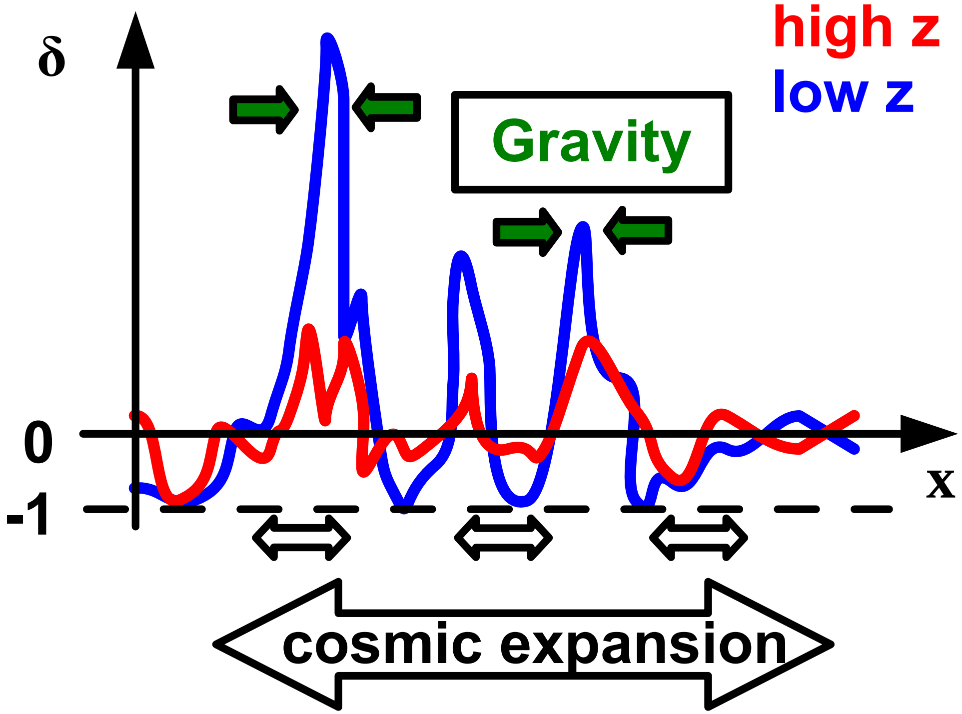
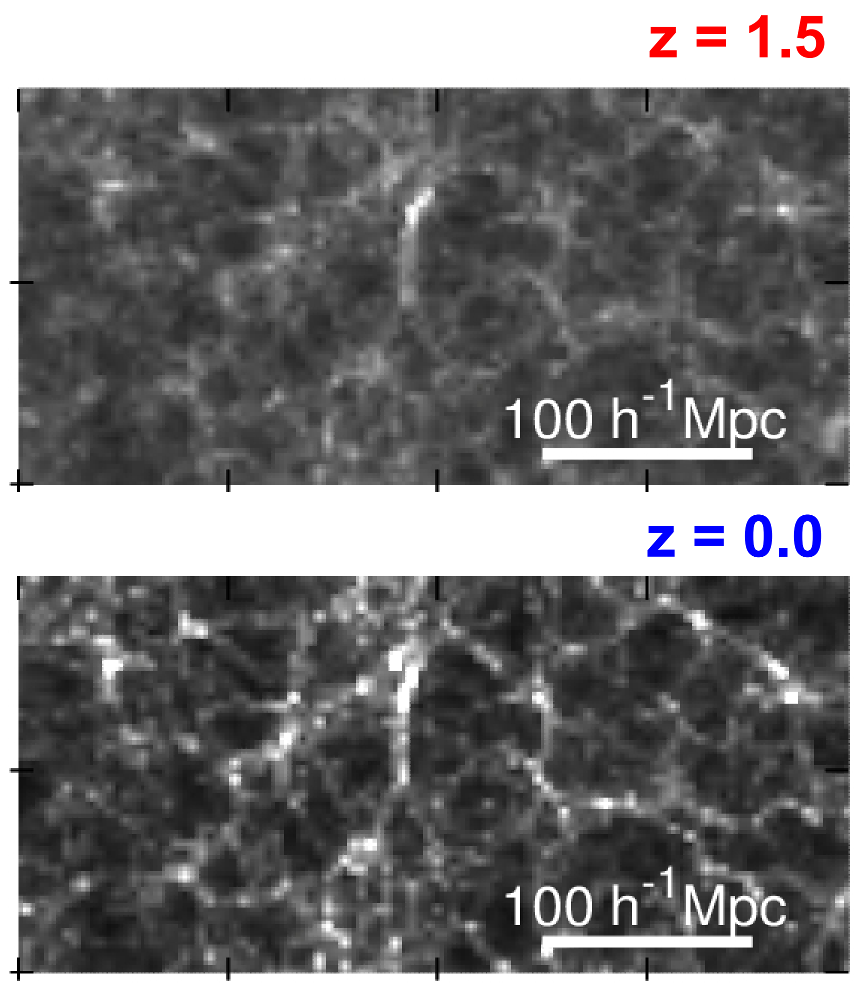
growth of matter fluctuations

matter fluctuations

$$\delta(x) \equiv (\rho(x) - \bar{\rho})/\bar{\rho}$$





MICE Grand Challenge simulation

growth depends on cosmology

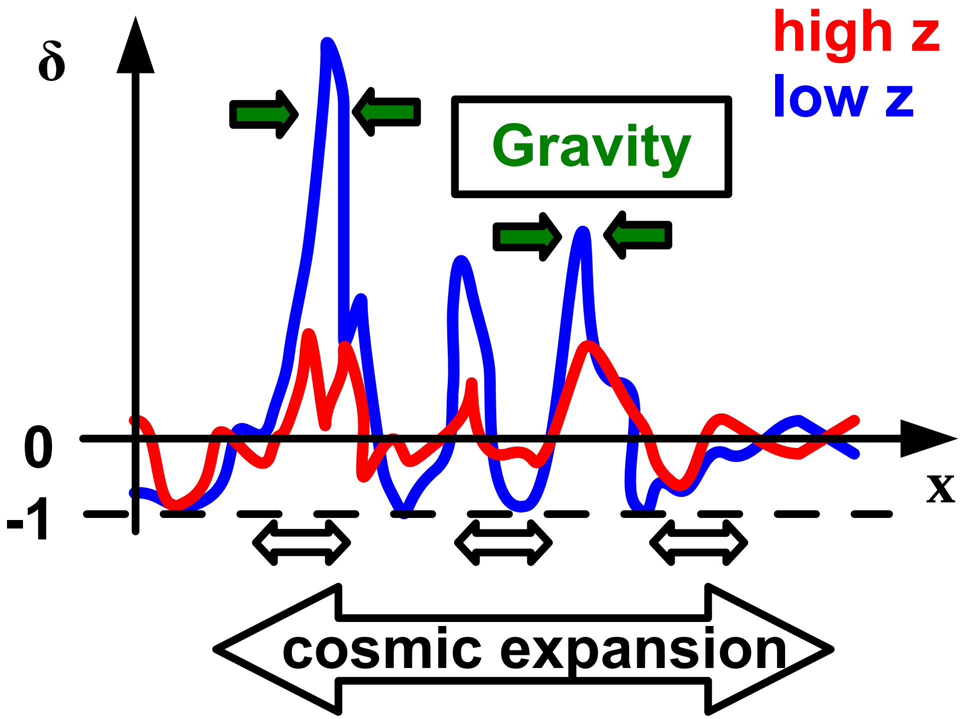
matter fluctuations

$$\delta(x) \equiv (\rho(x) - \bar{\rho}) / \bar{\rho}$$

linear growth factor

$$D(z) \simeq \delta_m(z) / \delta_m(z_0)$$

(large scale approximation)



growth sensitive to:

- expansion
- gravity
- matter density
- particle characteristics

two-point correlation depends on growth

two-point correlation (2pc)

$$\xi(r_{12}) \equiv \langle \delta_1 \delta_2 \rangle (r_{12})$$

linear growth factor

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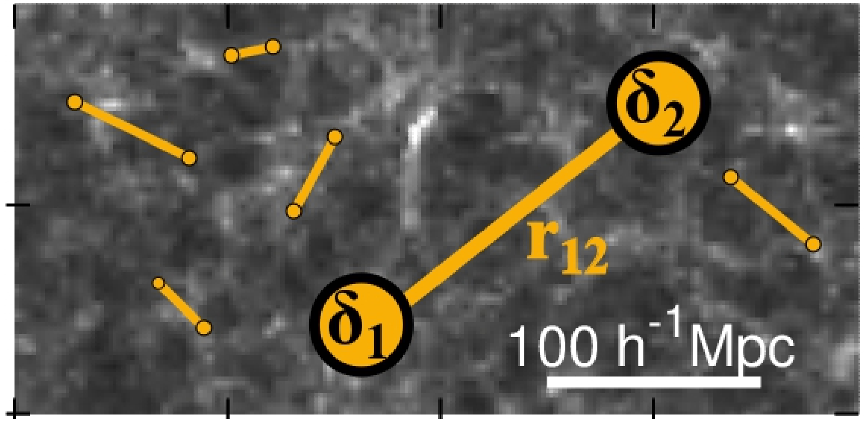

- study in 3D config. space
- ξ is isotropic

cosmology with large-scale structure

two-point correlation (2pc)

$$\xi(r_{12}) \equiv \langle \delta_1 \delta_2 \rangle (r_{12})$$

linear growth factor

$$D(z) \simeq \delta_m(z)/\delta_m(z_0)$$

(large scale approximation)

$$D_0^2(z) \simeq \xi_m(z)/\xi_m(z_0)$$

2pc growth measurement

compare with theory predictions, e.g. ΛCDM

(constrain free parameters, discriminate models)

$$D(a) \propto \frac{H(t)}{H(0)} \int_0^a \frac{da'}{\left[\Omega_m/a' + \Omega_\Lambda a' - (\Omega_m + \Omega_\Lambda - 1)\right]^{3/2}} \qquad a \equiv 1/(1+z)$$

cosmology with large-scale structure

two-point correlation (2pc)

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linear growth factor

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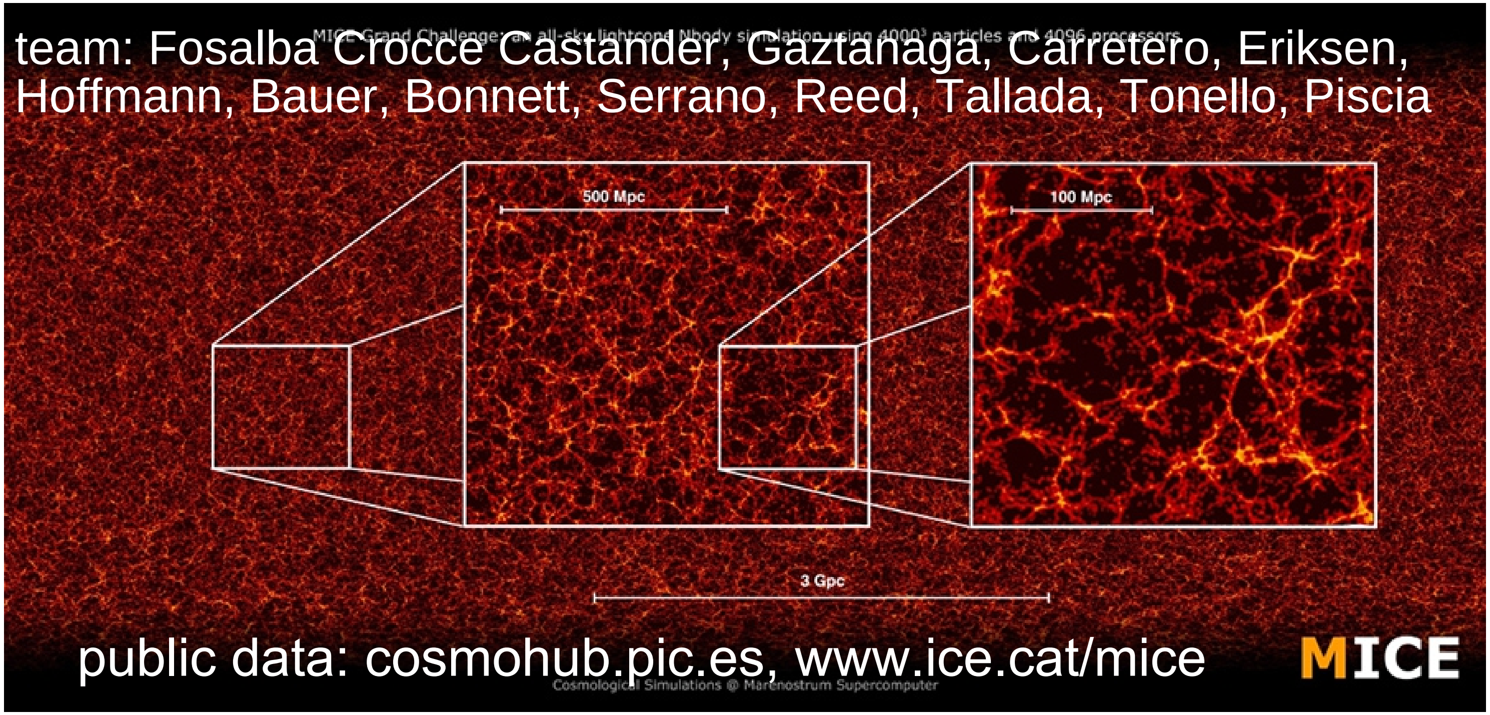
(large scale approximation)

$$D_0^2(z) \simeq \xi_m(z) / \xi_m(z_0)$$

2pc growth measurement

problem: matter 2pc not directly observable!

MICE Grand Challenge simulation



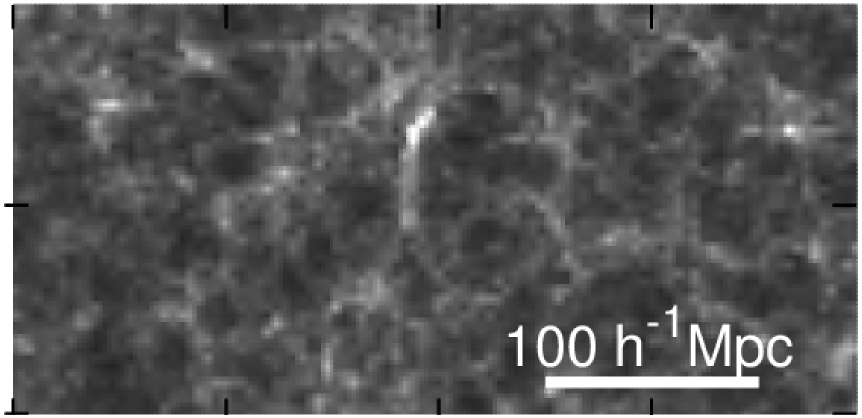
MareNostrum super computer

- 4096^3 ($7\ 10^{10}$) particles
- particle mass = $3\ 10^{10}$ Msun/h
- simulation box (3 Gpc/h)^3
- ΛCDM cosmology:

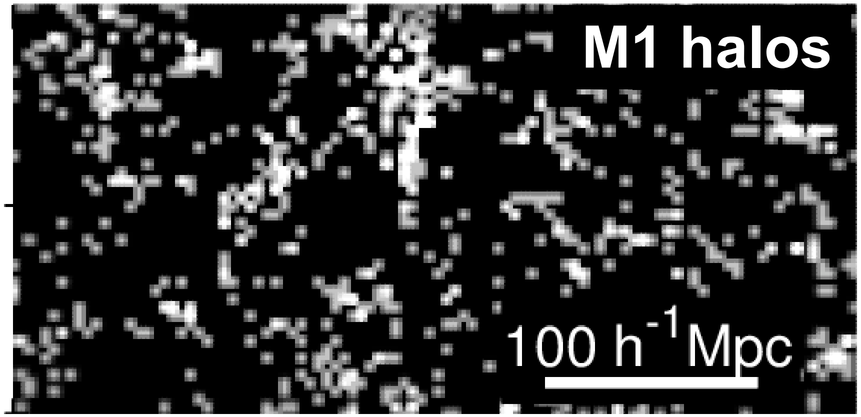
$$\Omega_m = 0.25, \Omega_\Lambda = 0.75 \Omega_b = 0.044, \sigma_8 = 0.8 n_s = 0.95, h = 0.7$$

galaxies trace matter density field

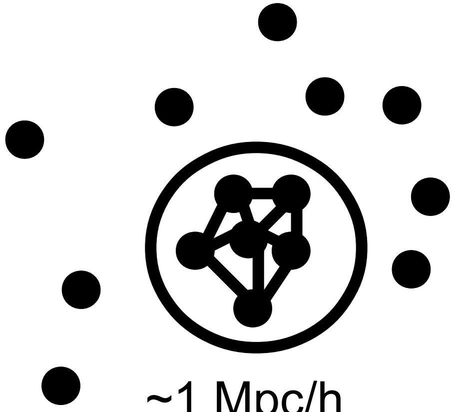
matter density



galaxy density



friends-of friends halo detection:

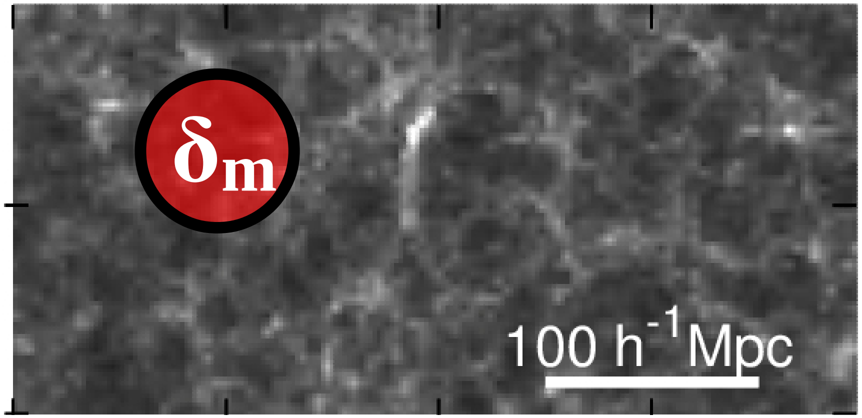


halo mass samples

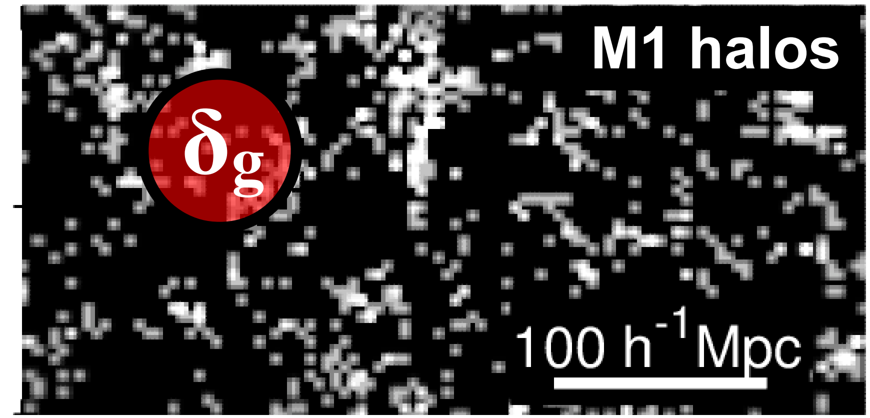
sample	mass range [10^{12} M_{sun}/h]
M0	0.58 - 2.32
M1	2.32 - 9.26
M2	9.26 - 100
M3	>100

galaxies bias model

matter density



galaxy density



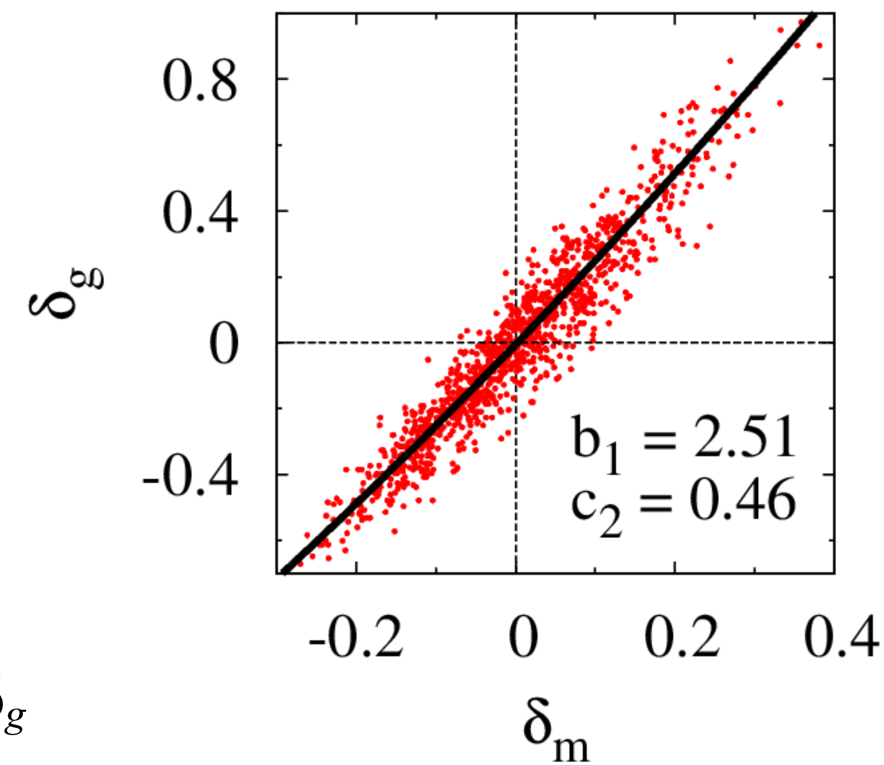
Fry & Gaztanaga (1993)

local quadratic bias model

$$\delta_g \simeq b_1 \left\{ \delta_m + (c_2/2)\left(\delta_m^2 - \langle \delta_m^2 \rangle\right) \right\}$$

- b_1: linear bias parameter
- c_2: quadratic bias parameter

assumption: δ_m determines δ_g



galaxies bias model

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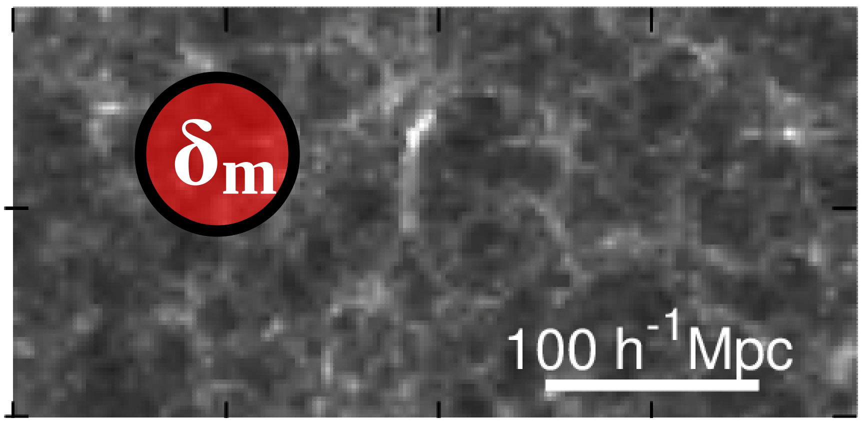
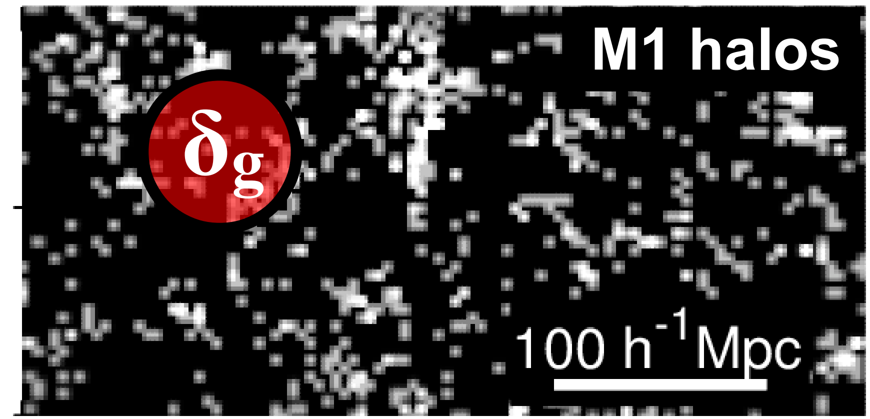


galaxy density



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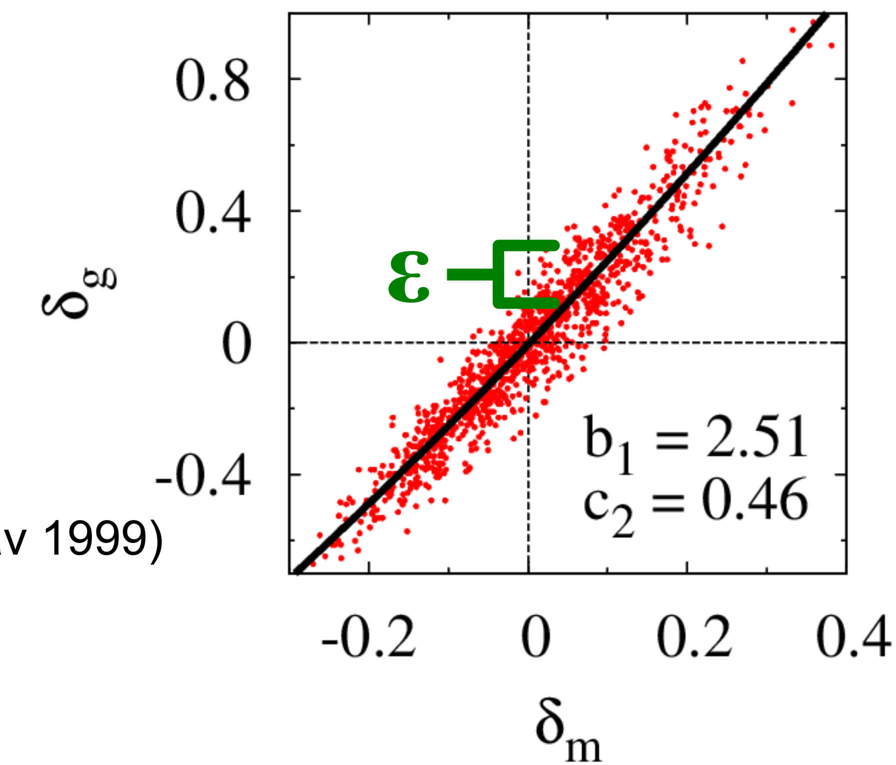
local quadratic bias model

$$\delta_g \simeq b_1 \left\{ \delta_m + (c_2/2)\left(\delta_m^2 - \langle \delta_m^2 \rangle\right) \right\}$$

+ ε **residual** from
  - stochasticity (e.g. Dekel, Lahav 1999)
  - "non-local" contributions (Chan 2012, Baldauf 2012)



bias in galaxy 2-point correlations

local quadratic bias model

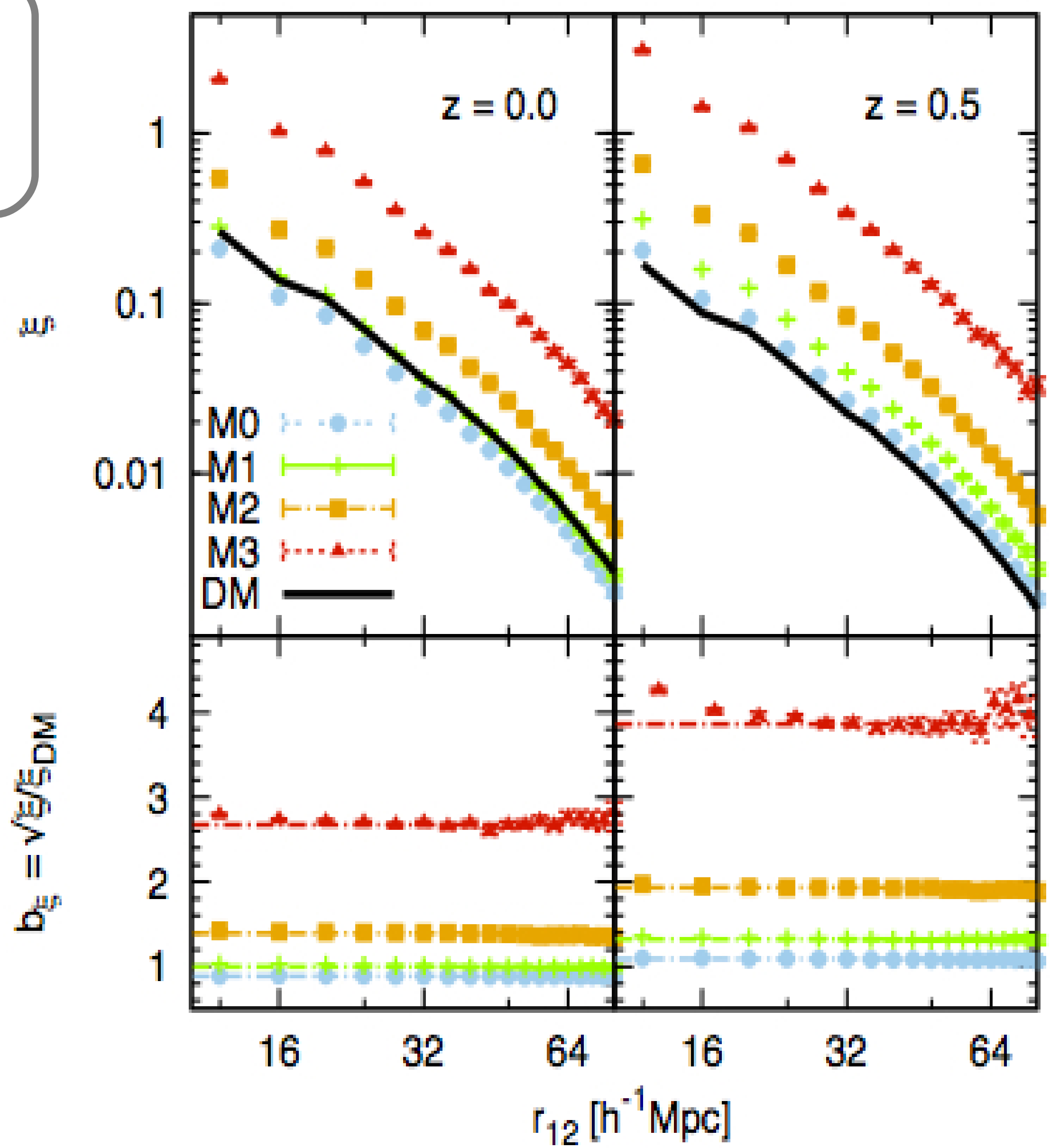
$$\delta_g \simeq b_1 \left\{ \delta_m + (c_2/2)\left(\delta_m^2 - \langle \delta_m^2 \rangle\right) \right\}$$

2-point correlation

$$\xi(r_{12}) \equiv \langle \delta_1 \delta_2 \rangle (r_{12})$$

leading order approximation for large scales

$$\xi_g \simeq \boxed{b_1^2} \xi_m$$

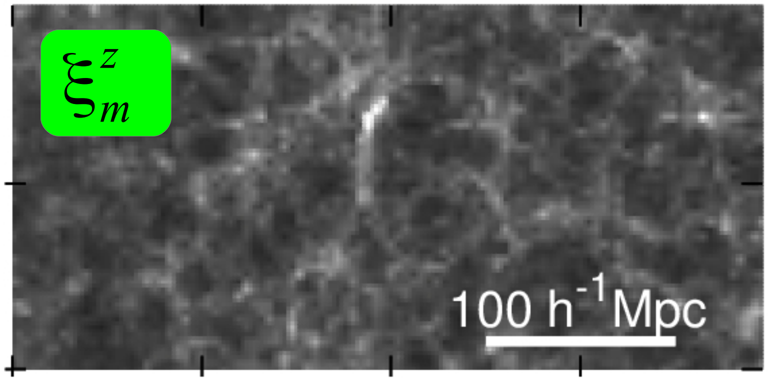


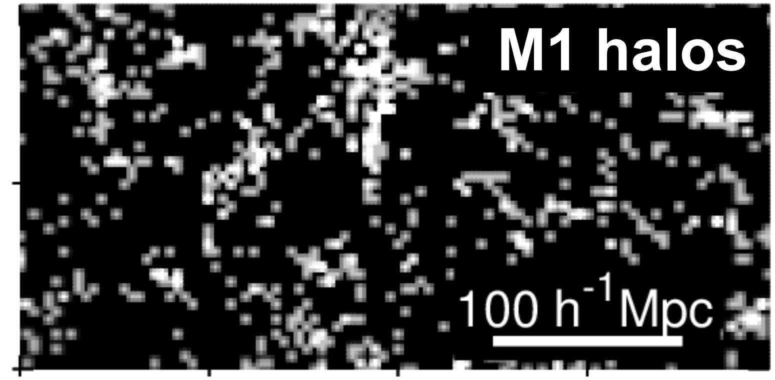
linear bias factor² (depends on redshift and halo mass)

growth – bias degeneracy

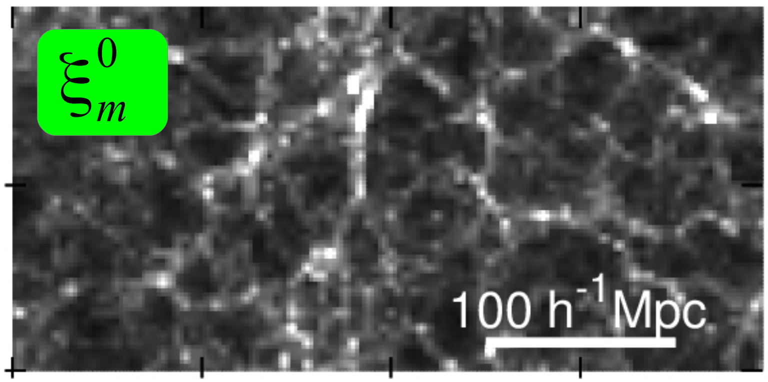
matter (theory) | **galaxies (observations)**

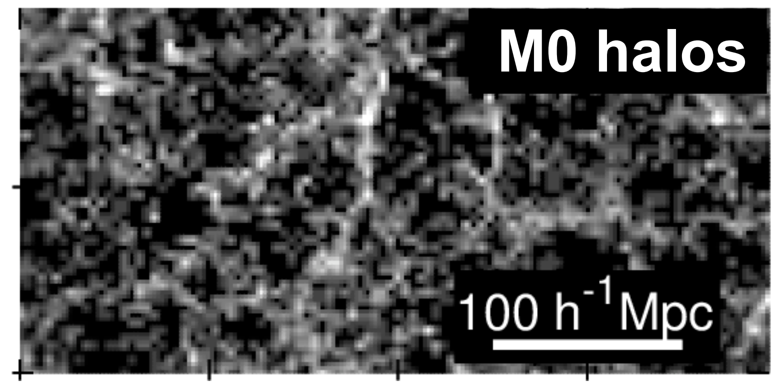
z = 1.0





z = 0.0





linear growth

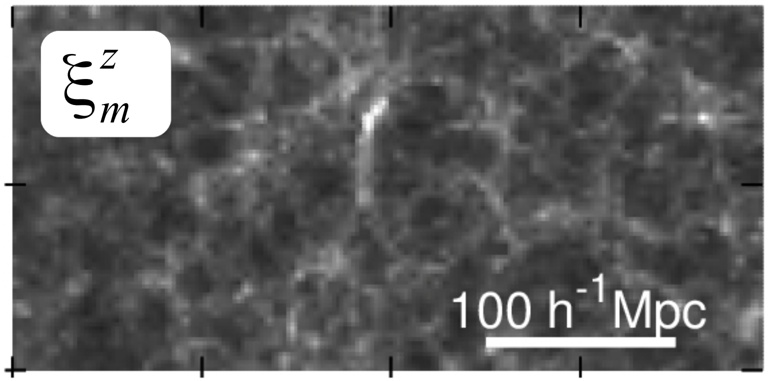
$$D_0(z)=\sqrt{\frac{\xi_m(z)}{\xi_m(0)}}$$

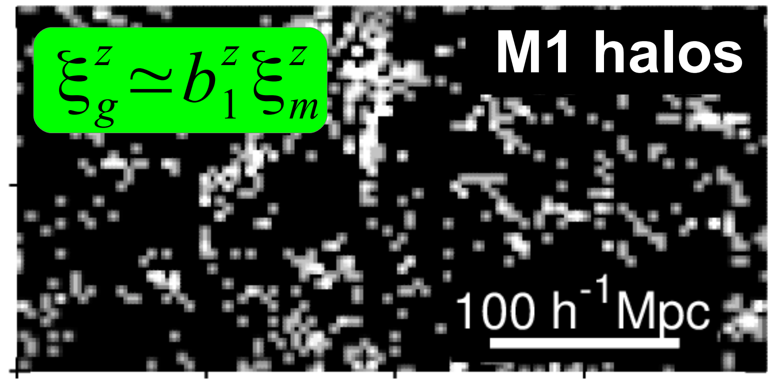
what we would like to observe

growth – bias degeneracy

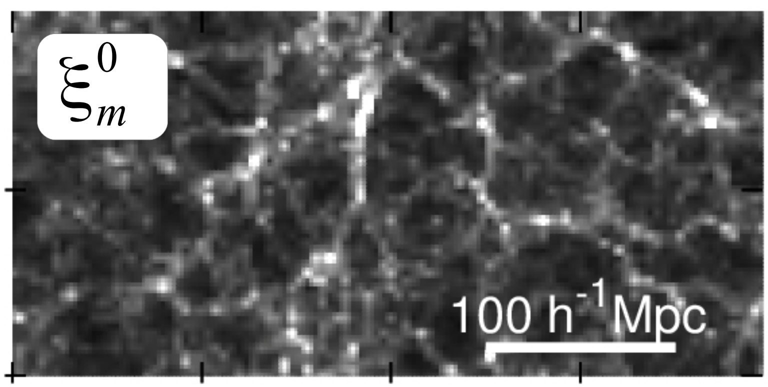
matter (theory) | **galaxies (observations)**

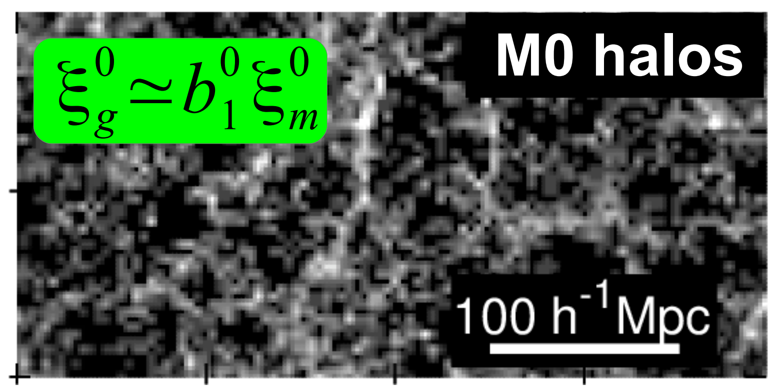
z = 1.0





z = 0.0





linear growth

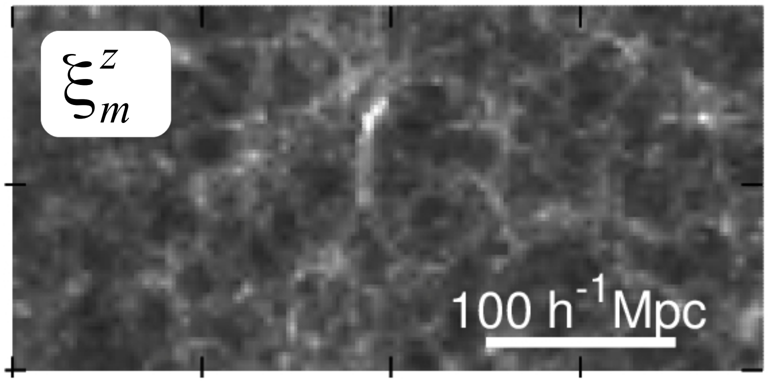
$$D_0(z) = \sqrt{\frac{\xi_m(z)}{\xi_m(0)}} = \frac{b(0)}{b(z)} \sqrt{\boxed{\frac{\xi_g(z)}{\xi_g(0)}}}$$

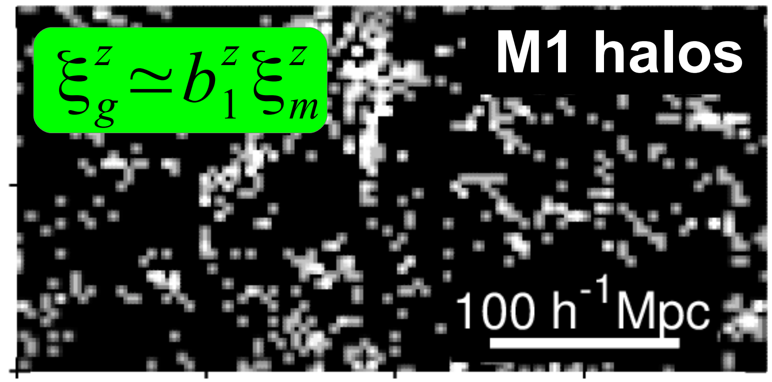
what we can observe

growth – bias degeneracy

matter (theory) | **galaxies (observations)**

z = 1.0





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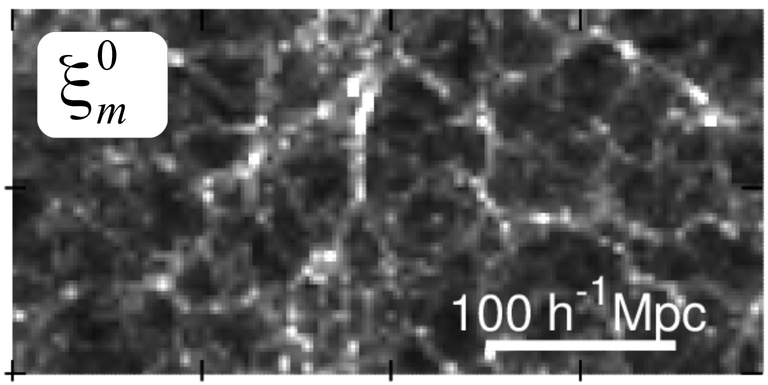
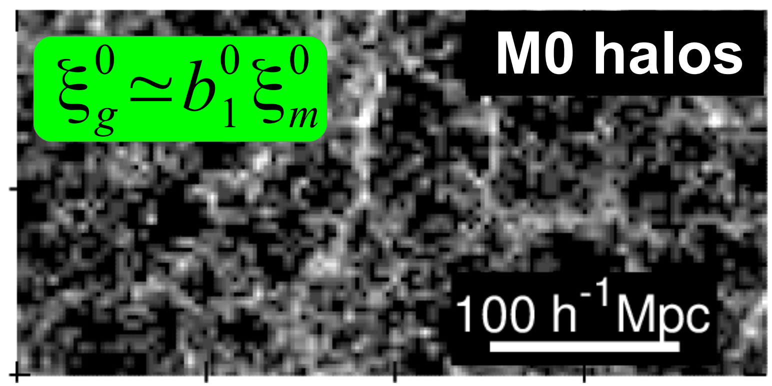




growth-bias degeneracy

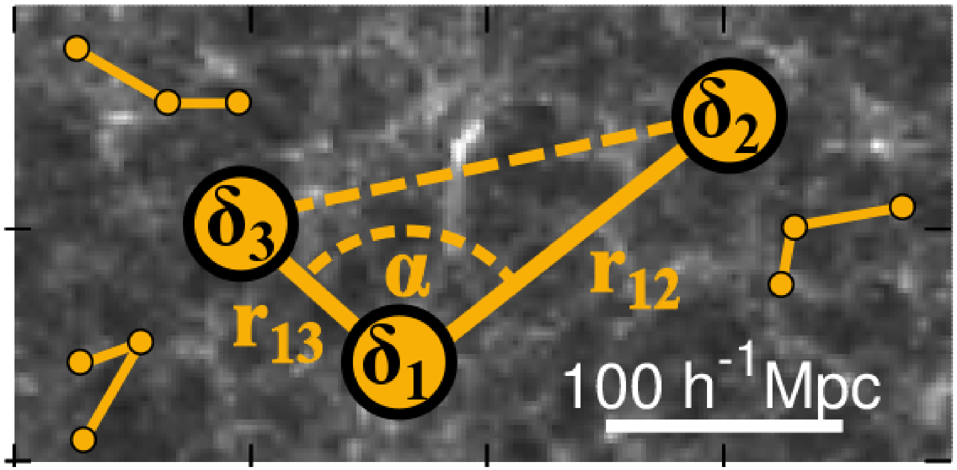
$$\boxed{D_0(z)\frac{b(z)}{b(0)}} = \sqrt{\frac{\xi_g(z)}{\xi_g(0)}}$$

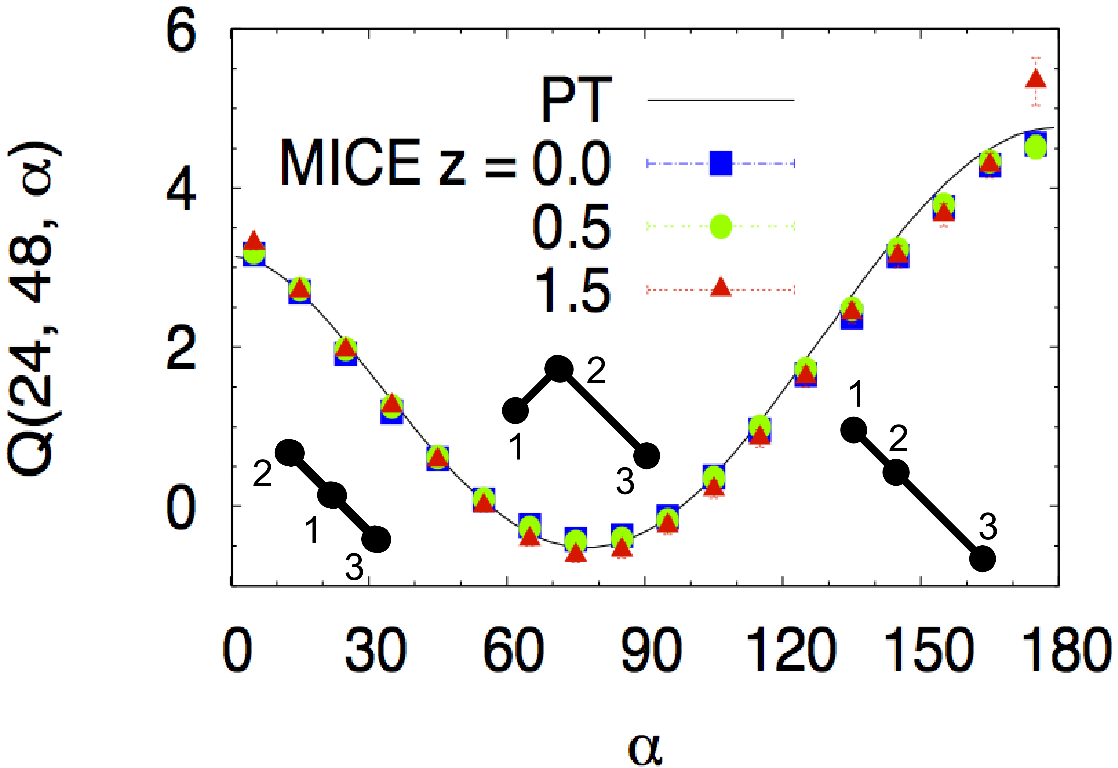
3-point correlation Q

reduced 3-point correlation:

$$Q \equiv \frac{\langle \delta_1 \delta_2 \delta_3 \rangle (r_{12}, r_{13}, \alpha)}{\langle \delta_1 \delta_2 \rangle \langle \delta_1 \delta_3 \rangle + 2\, perm.}$$

- **probes shape of LSS**





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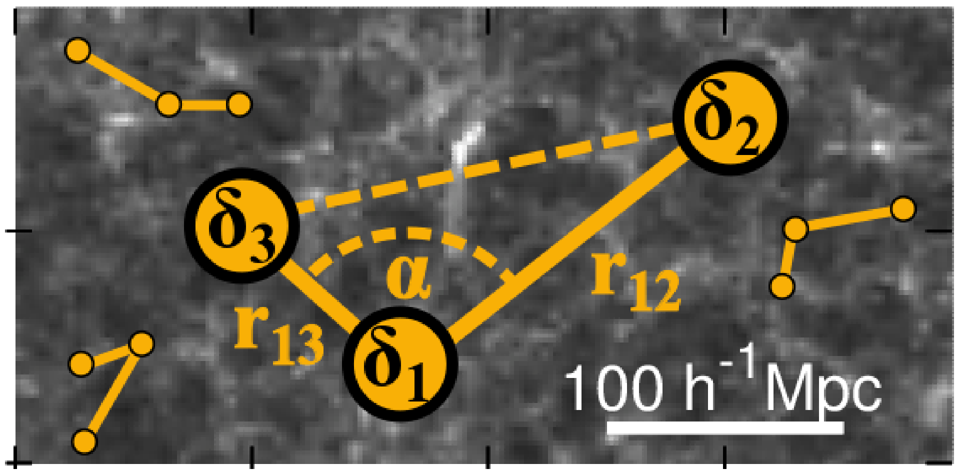
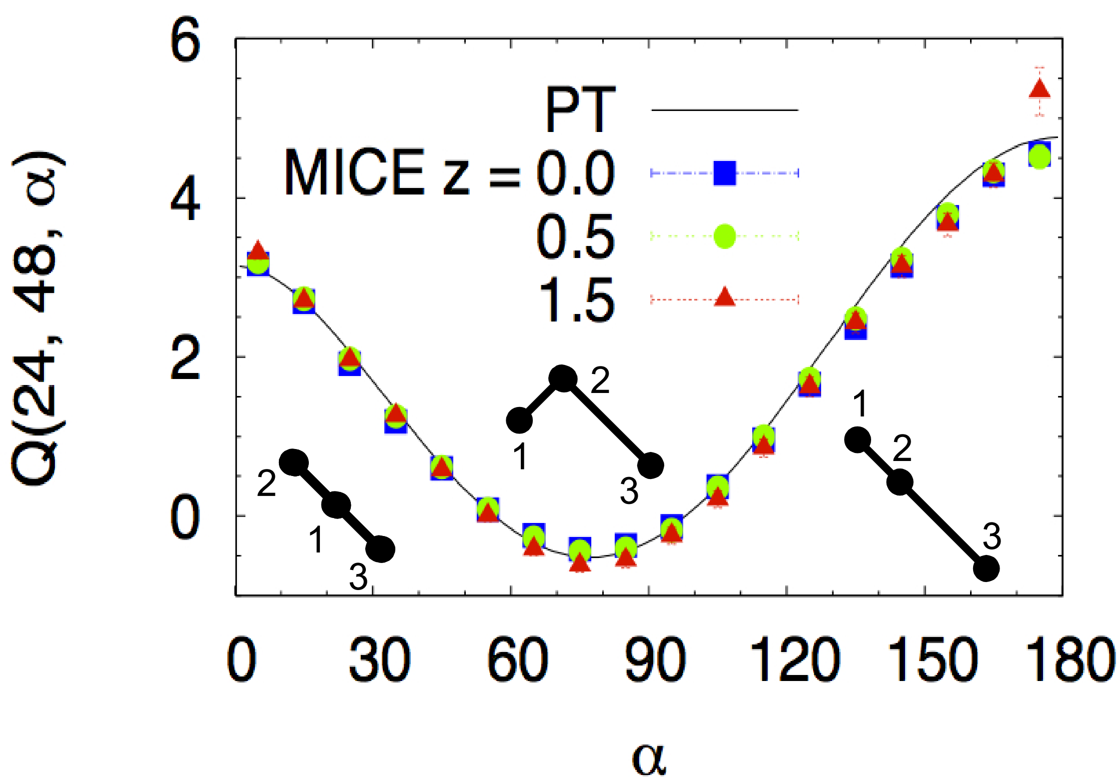
- probes shape of LSS
- **independent of redshift**





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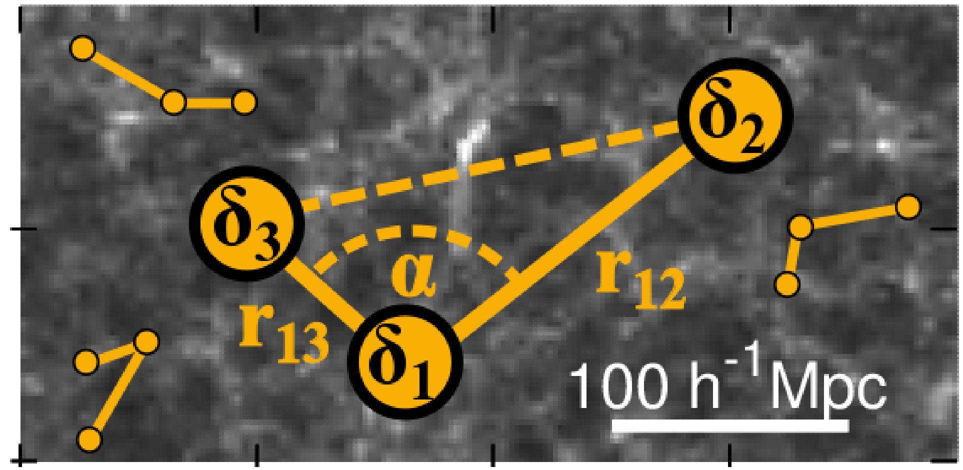
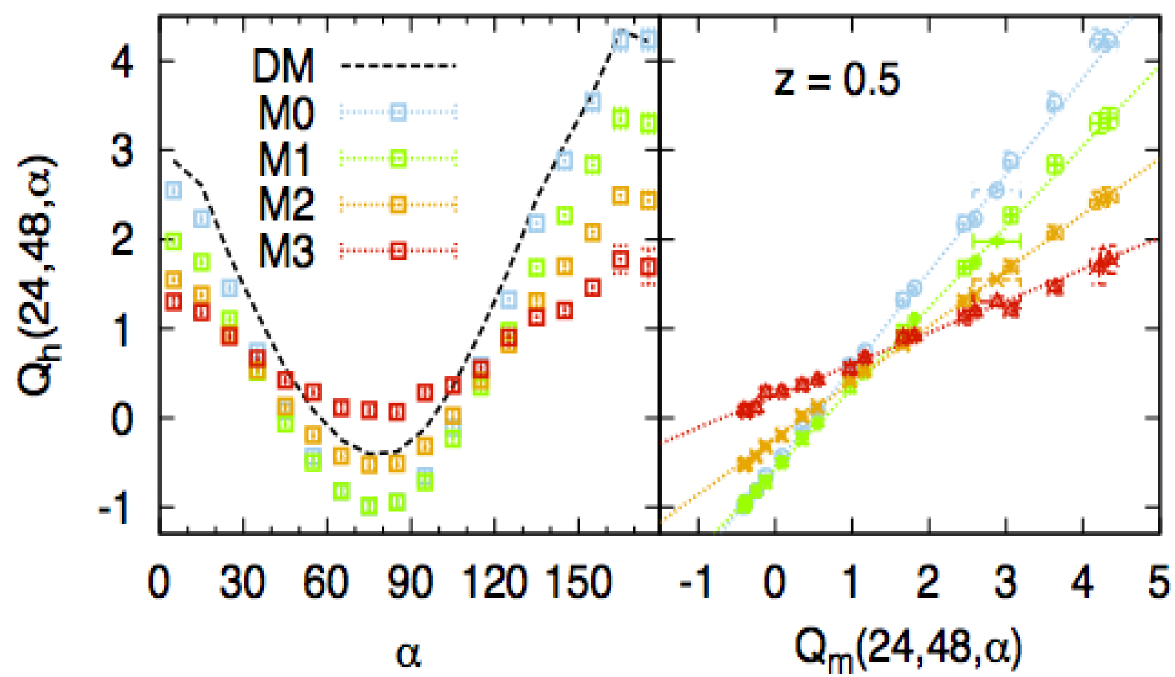


- probes shape of LSS
- independent of redshift
- **depends on bias**

$$\boxed{Q_g \simeq \frac{1}{b_1}(Q_m + c_2)}$$

- **large scale approximation**
- **based on local bias model**



Bel, Hoffmann, Gaztanaga 2015, arXiv:1403.1259

3-point correlation Q

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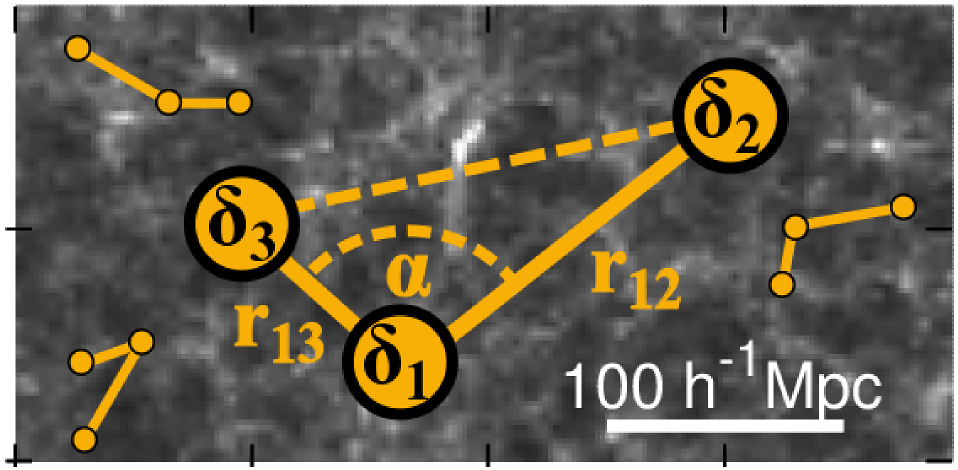
- probes shape of LSS
- independent of redshift
- depends on bias



growth - bias degeneracy broken with Q

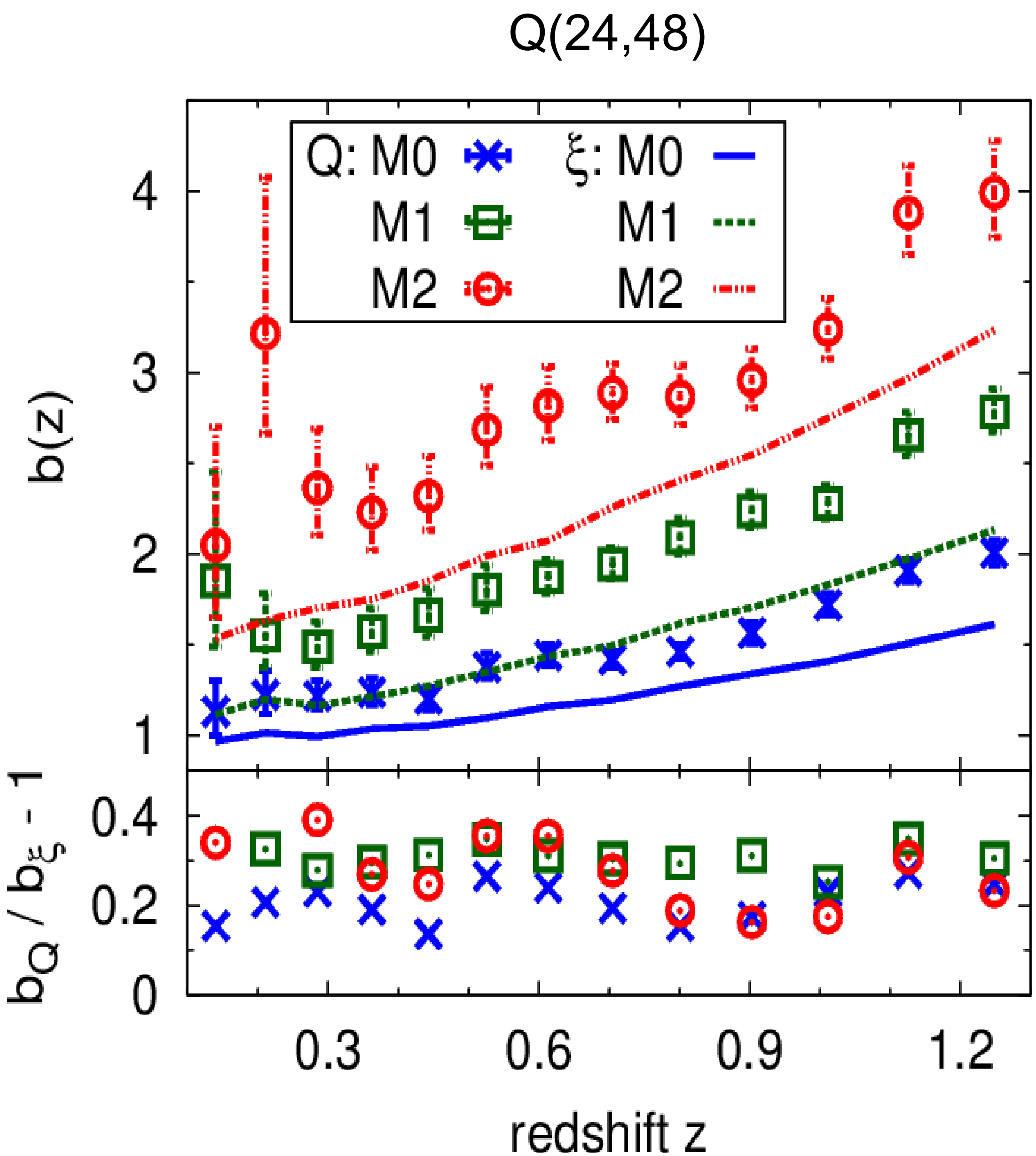
$$\boxed{Q_g \simeq \frac{1}{b_1}(Q_m + c_2)}$$

- **large scale approximation**
- **based on local bias model**

condition:

$$\frac{b_Q(0)}{b_Q(z)} = \frac{b_\xi(0)}{b_\xi(z)}$$

bias from Q ~30% too high



bias from ξ (lines)

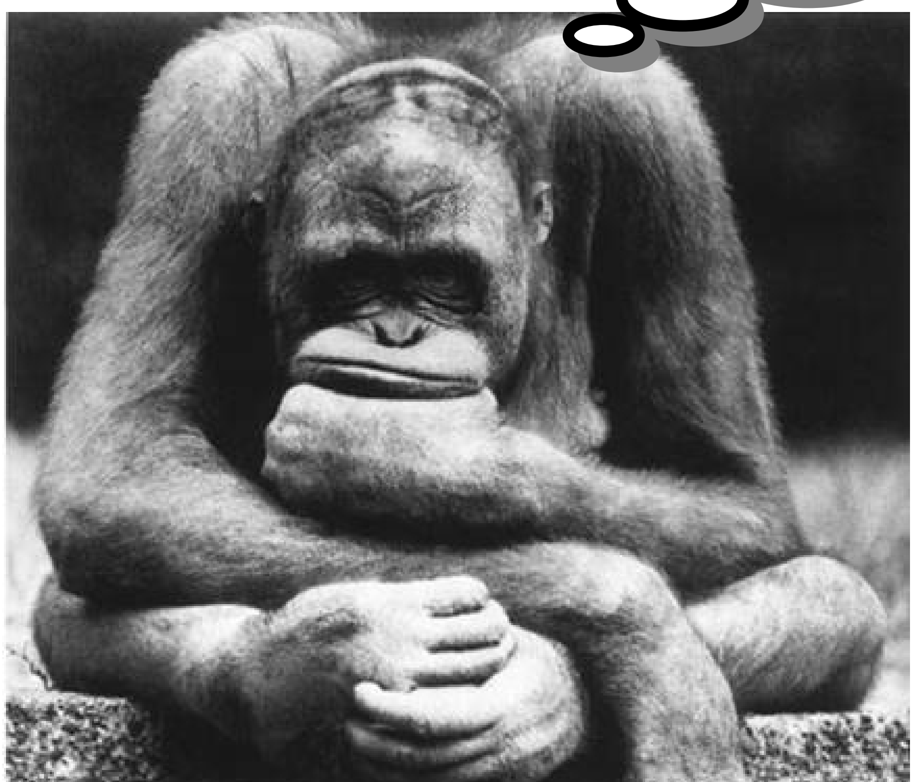
- good estimate of linear bias

bias from Q (symbols)

- **~30 % overestimation**
- overestimation similar for different halo mass samples and redshifts
- overestimation appears also very large scales, i.e. Q (36,72)

bias from Q ~30% too high

Why bias from Q too high?



- missing higher-order terms in Q_g ?
  (Pollack, Smith & Porciani, 2012)

- non-local bias ?
  (Chan, Scoccimarro & Sheth, 2012
  Baldauf et al. 2012)

non-local bias

local quadratic bias model:

$$\delta_g \simeq b_1 \delta_m + (b_2/2)(\delta_m^2 - \langle \delta_m^2 \rangle)$$

non-local bias

non-local quadratic bias model:

$$\delta_g \simeq b_1 \delta_m + (b_2/2)(\delta_m^2 - \langle \delta_m^2 \rangle)$$
$$+ \gamma_2 G_2(\Phi_v)$$

non-local contributions

$$G_2(\Phi_v) = (\nabla_{ij} \Phi_v)^2 - (\nabla^2 \Phi_v)^2$$

Chan, Scoccimarro & Sheth (2012)
Baldauf et al. 2012)

velocity potential: Φ_v

non-local bias: $\gamma_2 = g_2\, b_1/2$

non-local bias in 3-point correlation

non-local quadratic bias model

$$\delta_g \simeq b_1 \delta_m + (b_2/2)(\delta_m^2 - \langle \delta_m^2 \rangle)$$
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Q auto: galaxy-galaxy-galaxy

$$Q_g \simeq \frac{1}{b_1}(Q_m + [c_2 + g_2 Q_{nloc}])$$

non-local contributions

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non-local contributions

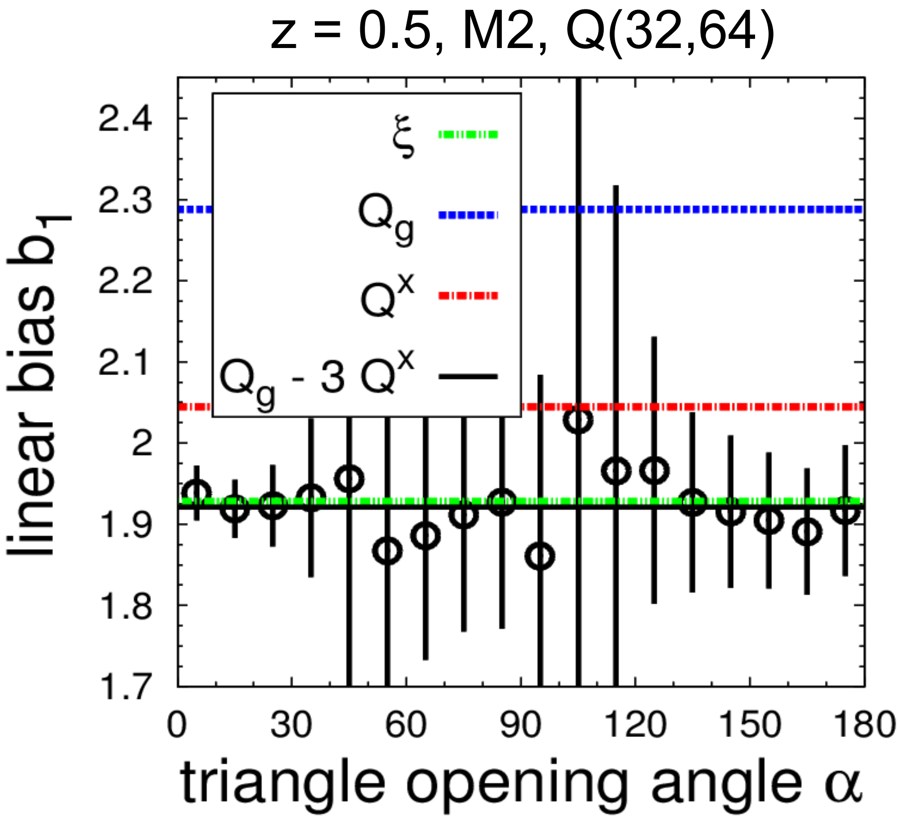
Q auto: galaxy-galaxy-galaxy

$$Q_g \simeq \frac{1}{b_1}(Q_m + [c_2 + g_2 Q_{nloc}])$$

Q cross: galaxy-matter-matter

$$Q^x \simeq \frac{1}{b_1}(Q_m + \frac{1}{3}[c_2 + g_2 Q_{nloc}])$$

non-local bias in 3-point correlation



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$$Q^x \simeq \frac{1}{b_1}\left(Q_m + \frac{1}{3}[c_2 + g_2 Q_{nloc}]\right)$$

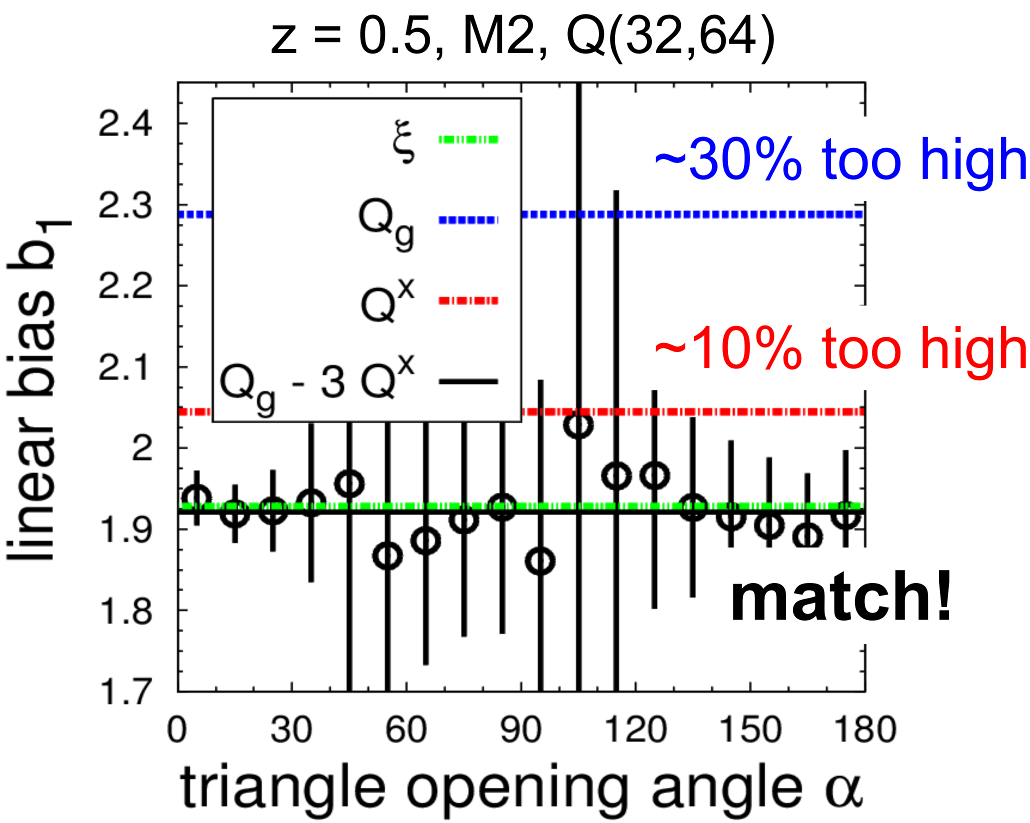
new linear bias estimator

Q auto – 3 Q cross:

$$b_1 = -2\,Q_m / \left(Q_h - 3\,Q^x\right)$$

- independent of quadratic and non-local contributions ($c_2 + g_2 Q_{nloc}$)
- **excellent match with "true" b_1 from ξ**
- possible application in galaxy-lensing cross-correlation

non-local bias in 3-point correlation



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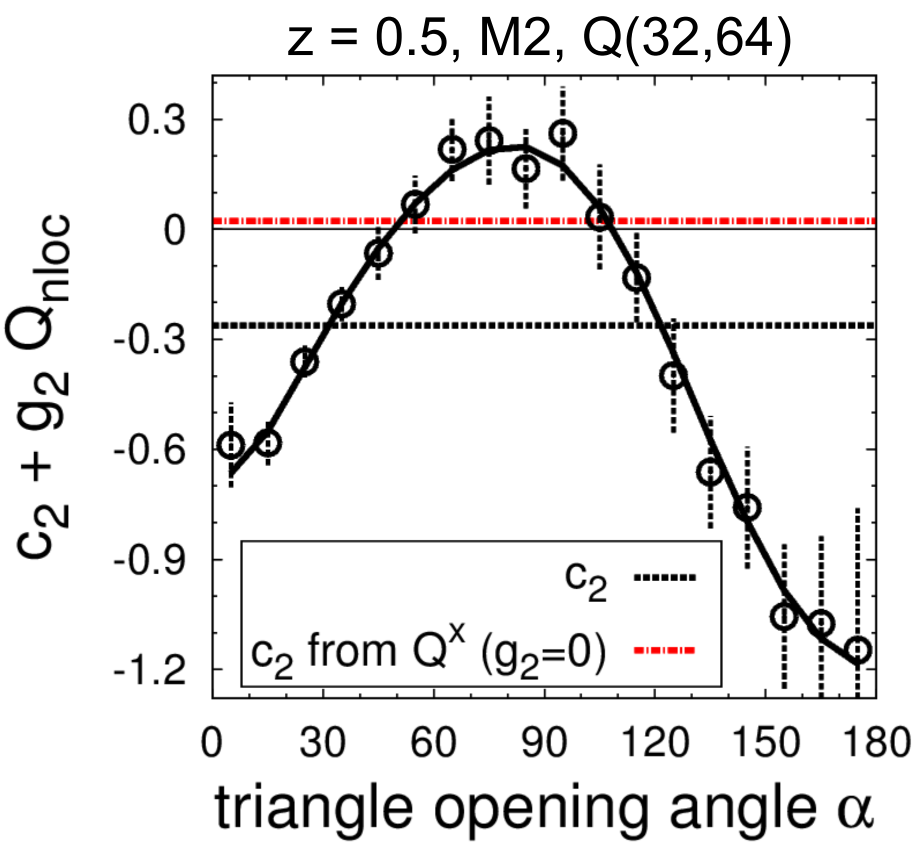

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$$Q_g \simeq \frac{1}{b_1}\left(Q_m + [c_2 + g_2 Q_{nloc}]\right)$$

non-local contributions

Q cross: galaxy-matter-matter

$$Q^x \simeq \frac{1}{b_1}\left(Q_m + \frac{1}{3}[c_2 + g_2 Q_{nloc}]\right)$$

Q auto – Q cross:

$$b_\xi(3/2)(Q_h - Q^x) = c_2 + g_2 Q_{nloc}$$

- independent of Q_m
- depends on triangle configuration
  → **local quadratic bias model (g_2 = 0) fails**
- agreement with Q_{nloc} prediction
- can be used to measure c_2 and g_2

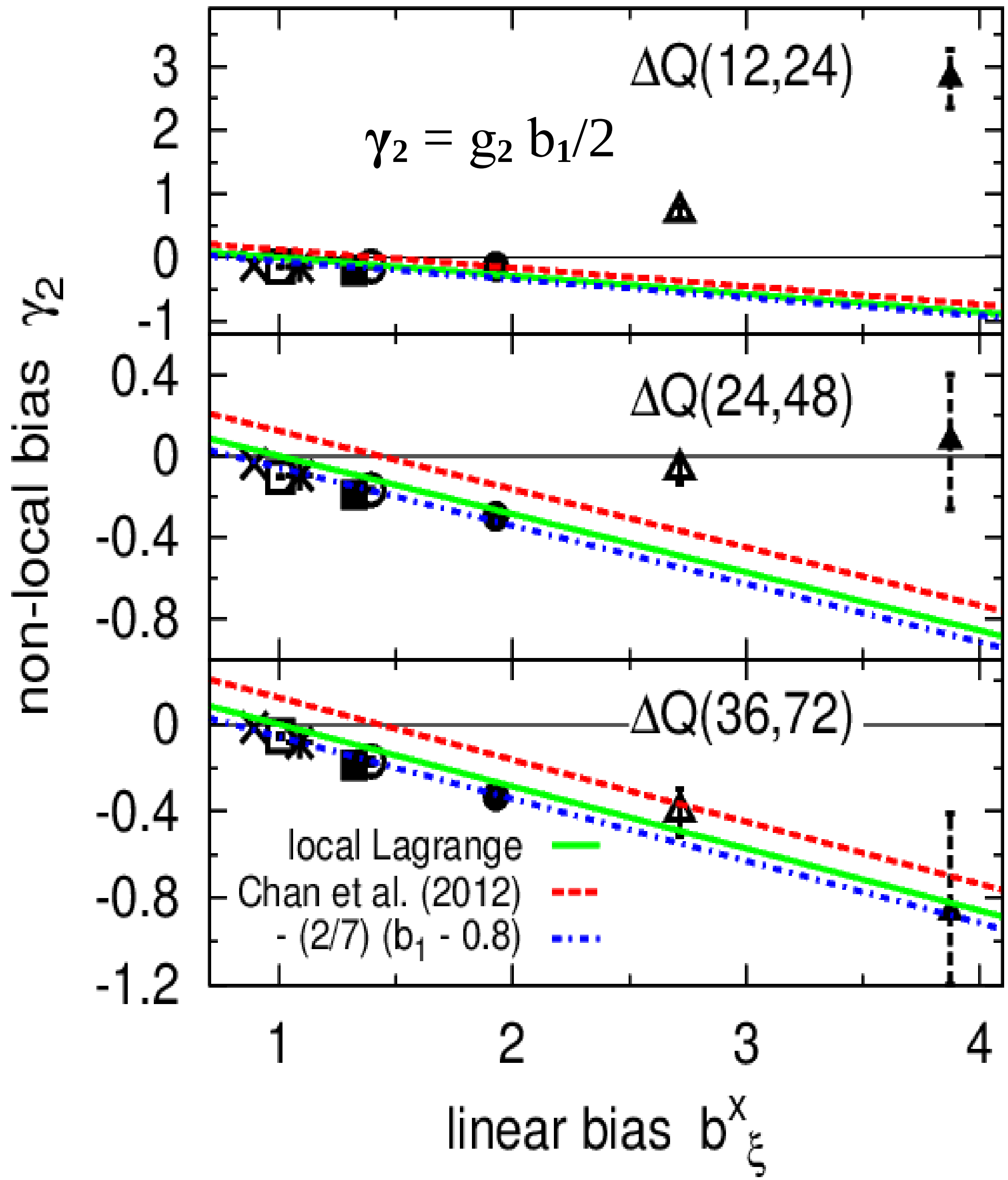
non-local bias in 3-point correlation

symbols:
- measurements from $\Delta Q = Q - Q^X$

lines:
- **fit to measurements**
- local Lagrangian prediction
- fit from Chan, Scoccimarro, Sheth (2012)

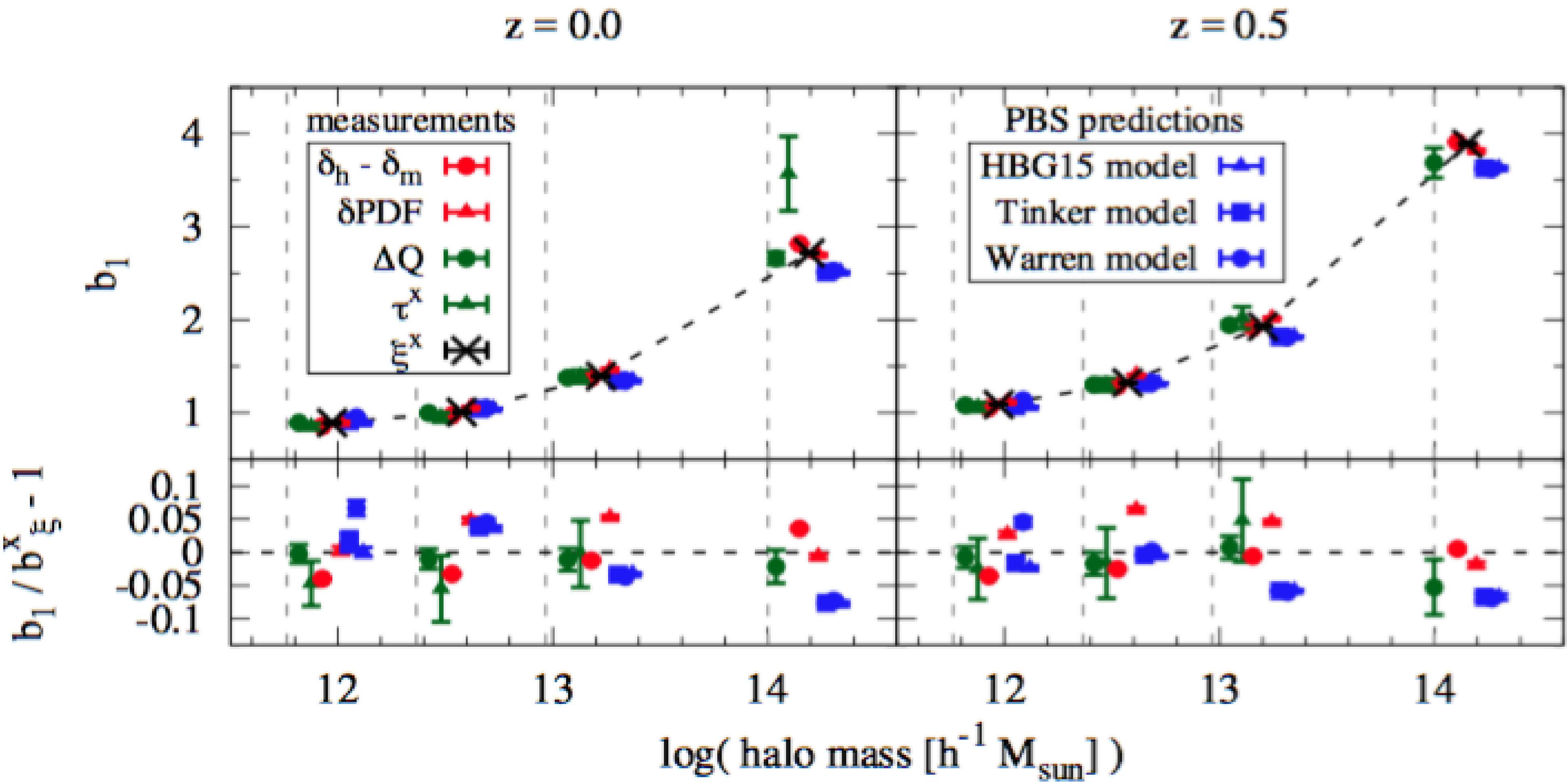
- **strong scale dependence of non-local bias $\gamma_2 = g_2\, b_1/2$**
- **at r > 35 Mpc/h**
  - **b_1-γ_2 linear**
  - **close to local lagrangian model**

$$g_2 = -(4/7)(1 - b_1^{-1})$$



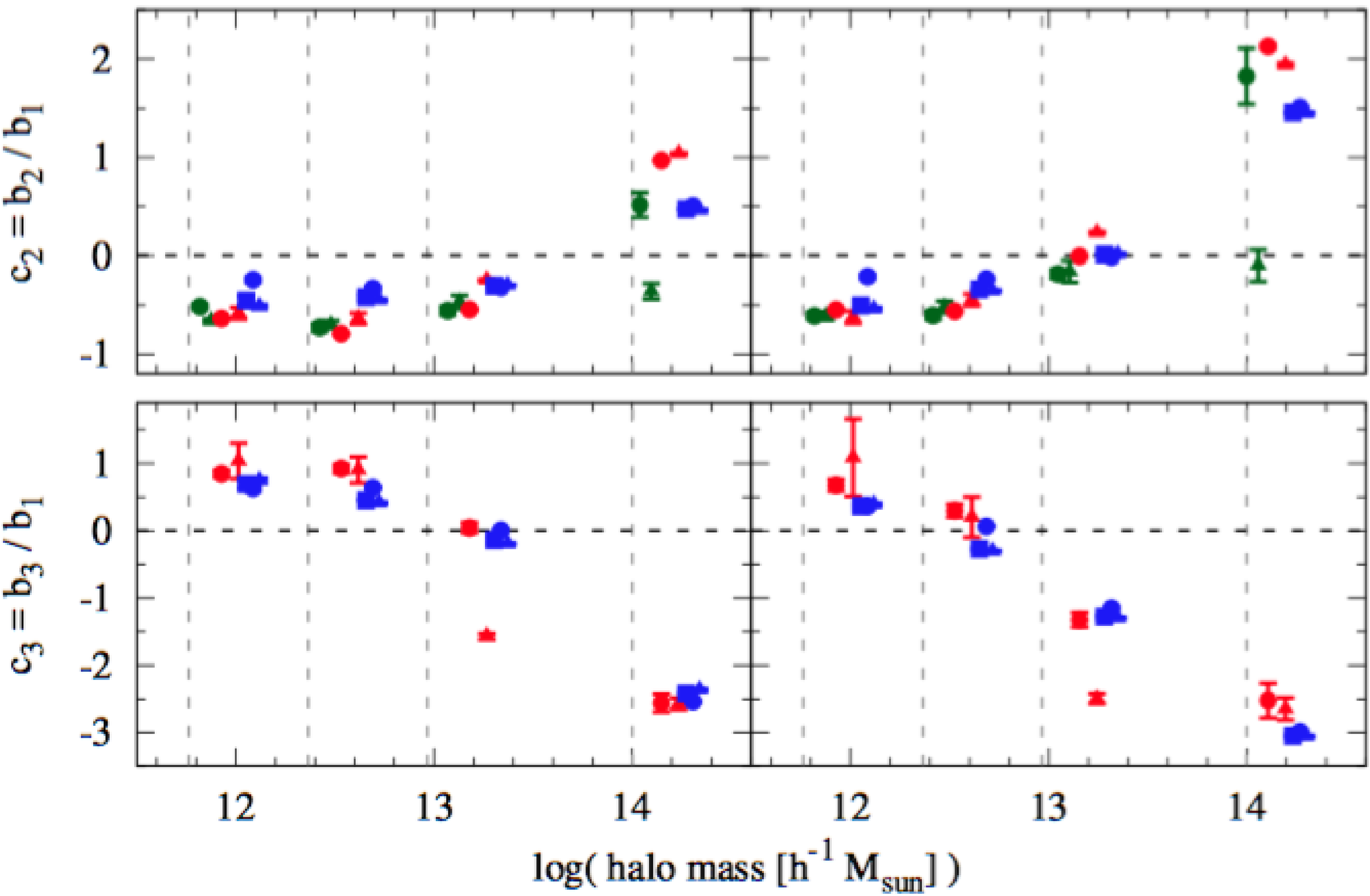
Hoffmann, Bel, Gaztanaga 2015, arXiv:1504.02074

linear bias comparison



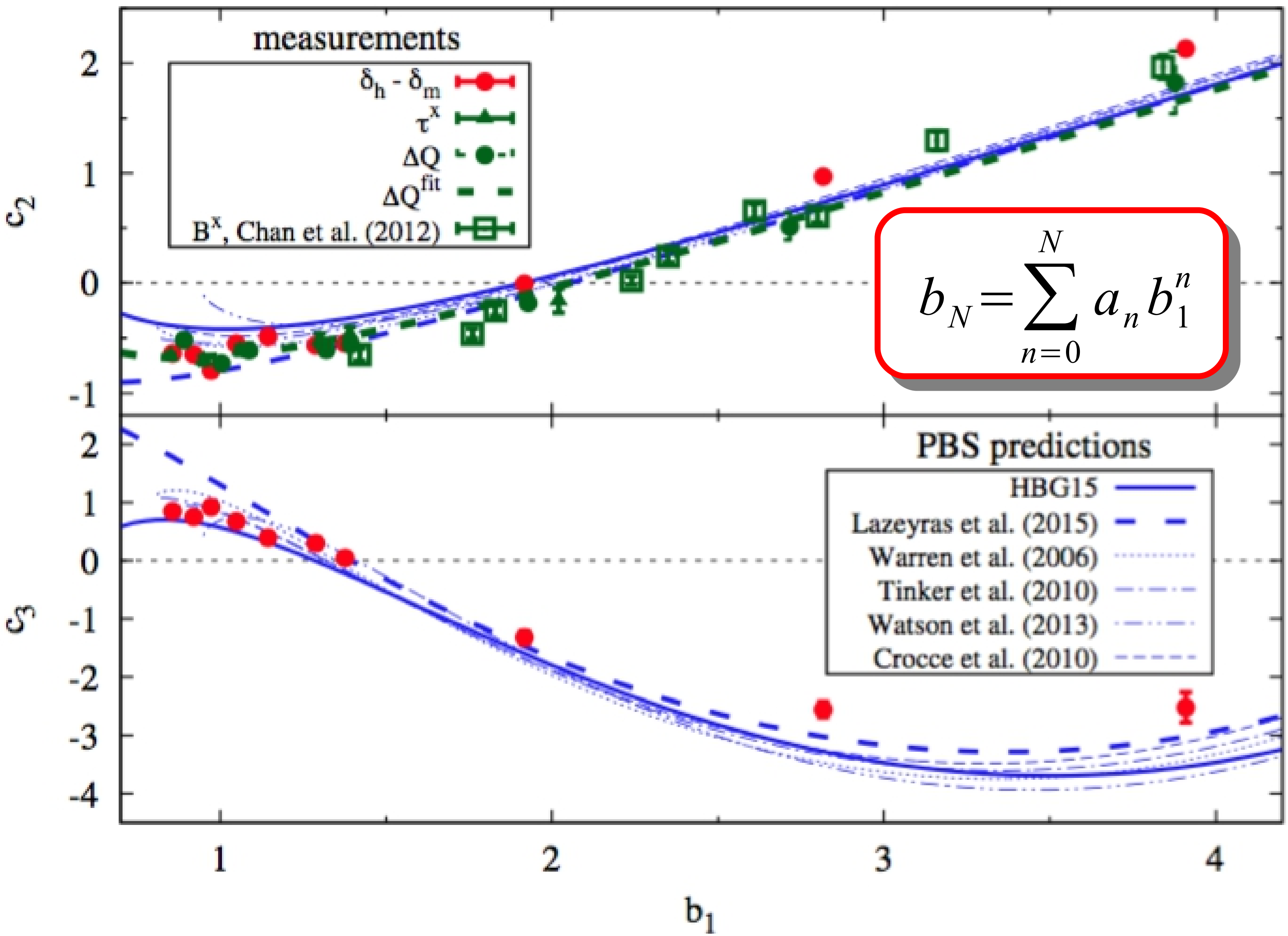
Hoffmann, Bel, Gaztanaga 2016, arXiv:1607.01024

non-linear bias comparison



Hoffmann, Bel, Gaztanaga 2016, arXiv:1607.01024

universal relation between b1, b2, b3



Hoffmann, Bel, Gaztanaga (2015, 2016, arXiv:1503.00313, 1607.01024)

b1 constraints from observed 3pc

$$\overbrace{Q_h}^{\text{measurement}} \simeq \frac{1}{b_1} \underbrace{\left(Q_m + [c_2 + g_2 Q_{nloc}]\right)}_{\text{model}}$$

model

- Qm & Qnloc from leading order perturbation theory
- c2 from (universal) c2(b1) relation *(Hoffmann, Bel, Gaztanaga 2015)*

  $$c_2 = 0.77 \mathrm{b}_1^{-1} - 2.43 + b_1$$

- g2 from local Lagrangian bias model

  $$g_2 = -(4/7)(1 - b_1^{-1})$$

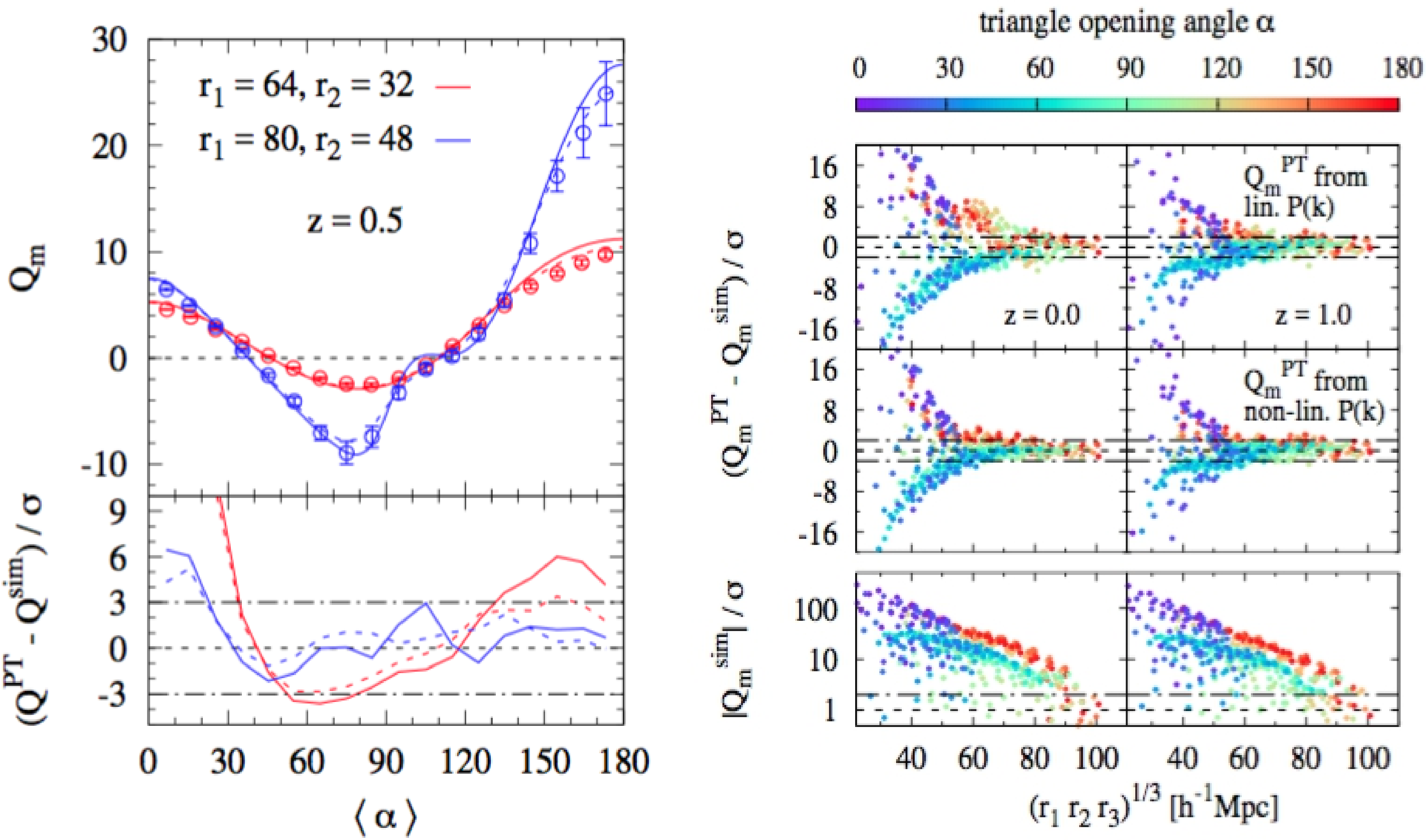
for measuring b1 from 3pc in surveys we would like to know:

- can b1 - c2 and b1-g2 relations be used to reduce
- number of free bias parameters (=> reduce error on b1) ?
- at which scales is PT valid?
- consistent results in config. and Fourier space ?

Hoffmann, Gaztanaga, Scoccimarro, Crocce in prep.:

- 49 N-body sims, 1280 (Mpc/h)^3, 3pc for 504 triangles

matter 3pc: PT vs simulation



- Qm predictions from leading order perturbation theory
- 2 sigma agreement with measurements for (r1r2r3)^1/3 > 60
- prediction improve when using non-lin Pk

halo 3pc: PT vs simulation

models:

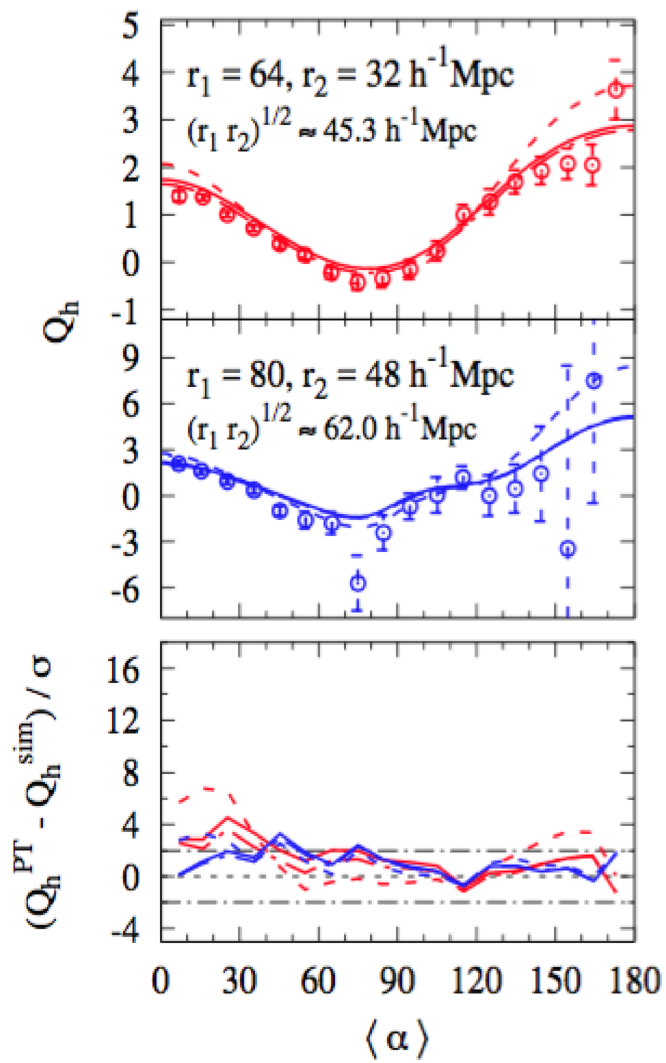
halo-halo-halo

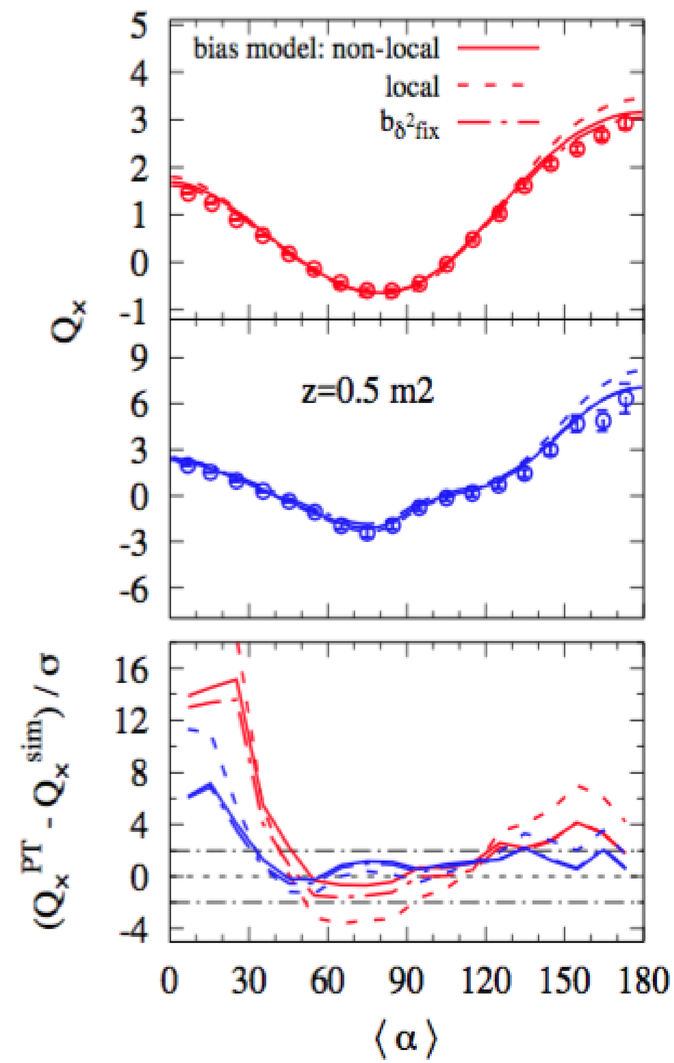
$$Q_h \simeq \frac{1}{b_1}\left(Q_m + [c_2 + g_2 Q_{nloc}]\right)$$

halo-matter-matter

$$Q_x \simeq \frac{1}{b_1}\left(Q_m + \frac{1}{3}[c_2 + g_2 Q_{nloc}]\right)$$

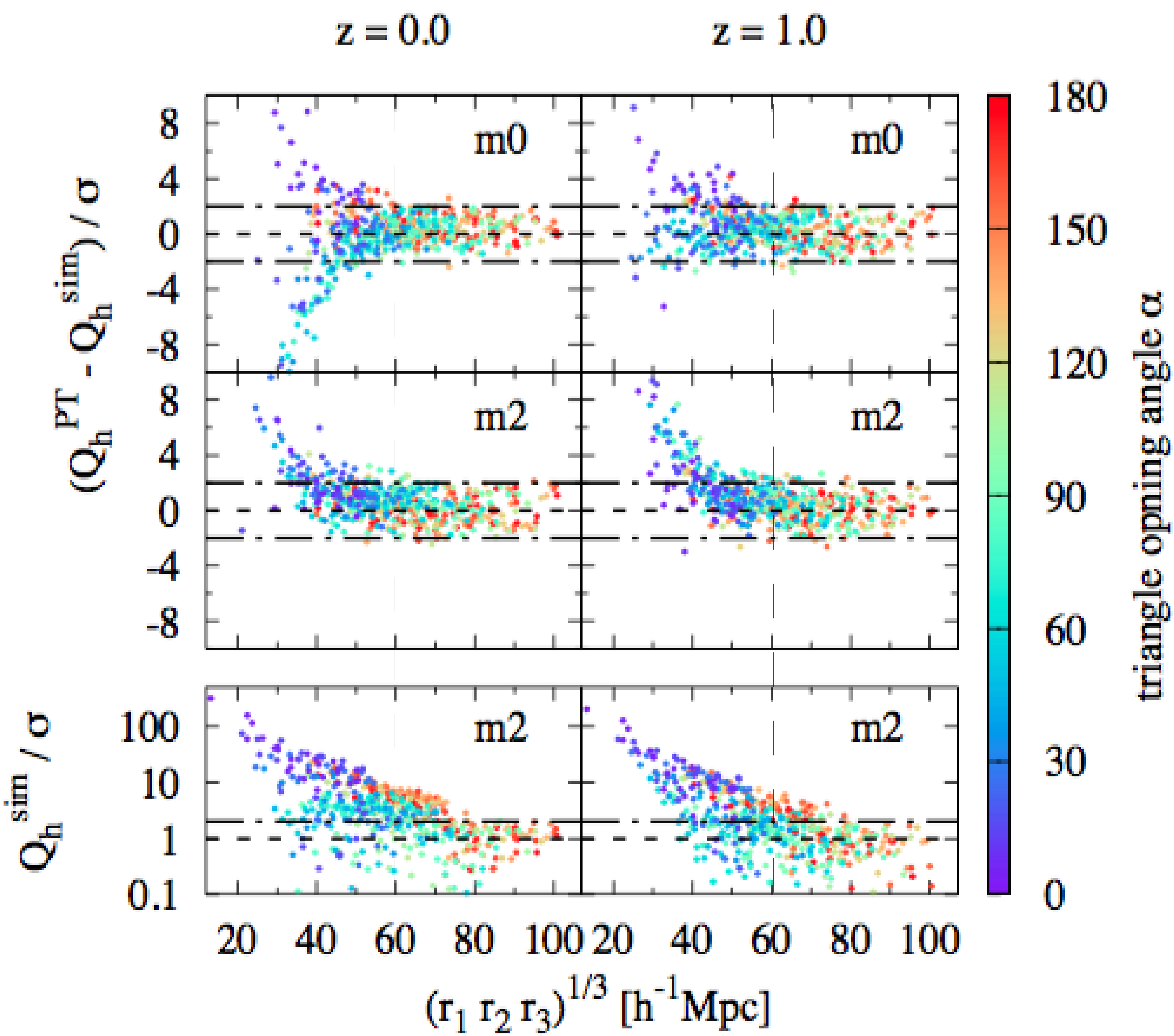
- Qm & Qnloc from leading order PT
- b1, c2, g2 from Fourier space measurements (Chan, Sheth, Scoccimarro, 2012)
- local bias model (g2=0, dashed) fails
- c2(b1) and g2(b1) relations (dashed-dotted) work!





halo-halo-halo 3pc: PT vs simulation

significance of deviations



triangle scale

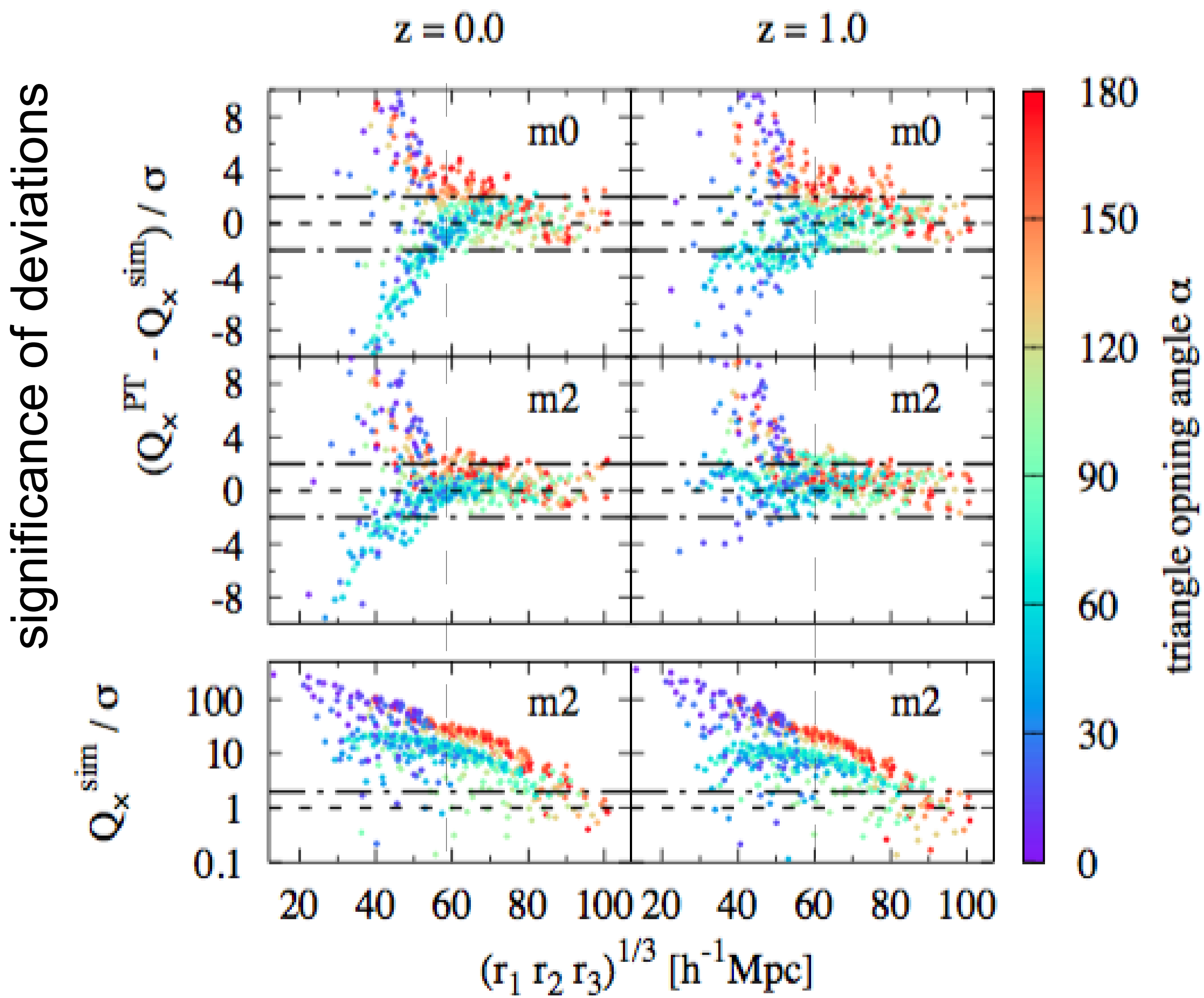
(r1 r2 r3)^1/3
> 60 Mpc/h:

- deviations < 2 sig
- consistent results in Fourier and config. space
- PT models for 3pc work

Conclusions

- constraints on cosmological models limited by growth–bias degeneracy
- growth–bias degeneracy broken with 3pc
- 3pc is affected by non-local contributions to bias function
- agreement with Fourier space results only at > 60 Mpc/h
- nearly universal relation between b2(b1) and b3(b1) can be used to decrease error on b1 from 3pc

halo-matter-matter 3pc: PT vs simulation



triangle scale

constraining lin bias (b1) with 3pc

non-local quadratic 3pc model

$$Q_h \simeq \frac{1}{b_1}\left(Q_m + [c_2 + g_2 Q_{nloc}]\right)$$

$$c_2 = 0.77 \mathrm{b}_1^{-1} - 2.43 + b_1$$

(roughly) universal b1 - c2 relation
(Hoffmann, Bel, Gaztanaga 2015)

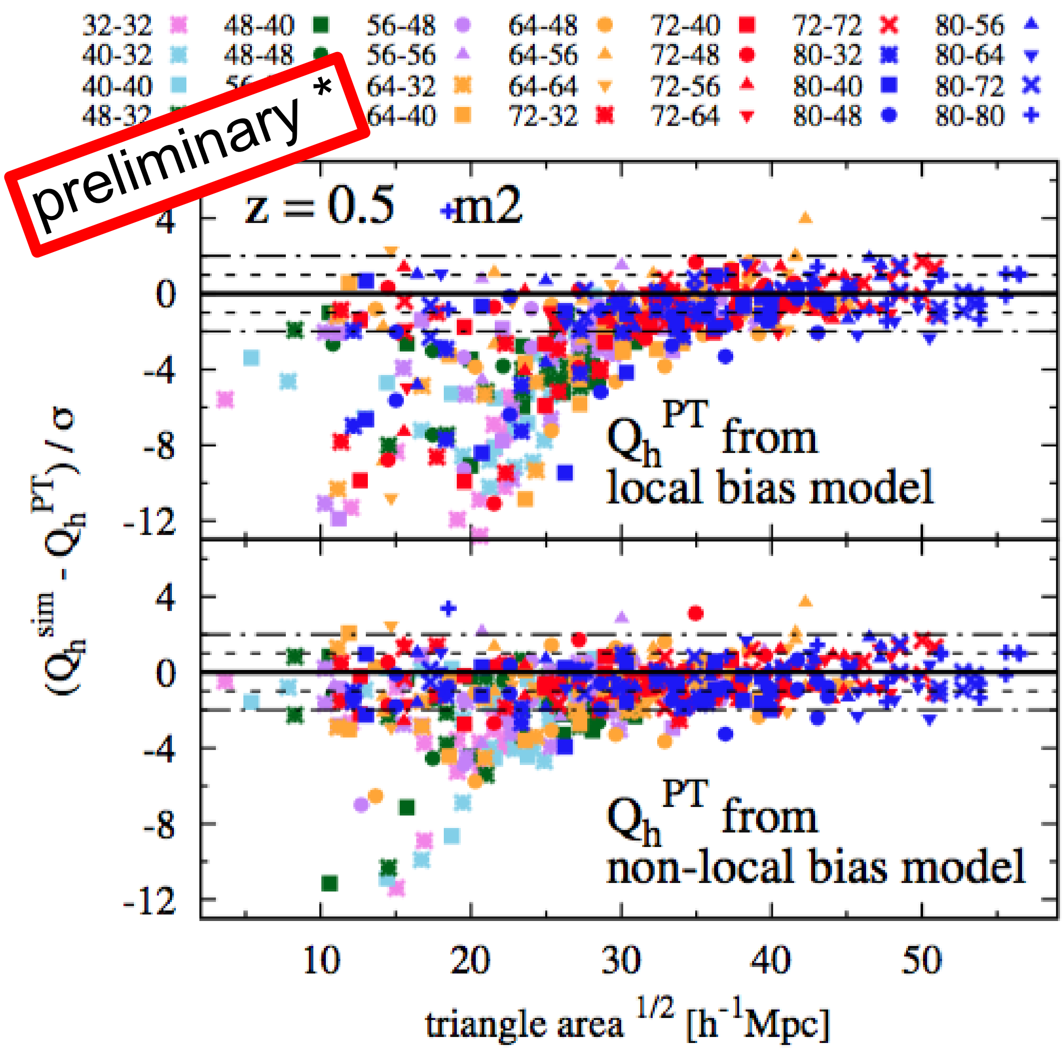
$$g_2 = -(4/7)\left(1 - b_1^{-1}\right)$$

local Lagrangian bias model

questions concerning an application in observations

- can b1 - c2 and b1-g2 relations be used to reduce Nr. of free bias parameters (=> reduce error on b1) ?
- at which scales is PT valid?

3pc in Fourier and config. space



$$\overbrace{Q_h}^{\text{measurement}} \simeq \frac{1}{b_1} \underbrace{(Q_m + [c_2 + g_2 Q_{nloc}])}_{\text{model}}$$

- Qm & Qnloc from leading order perturbation theory
- b1, c2, g2 from Fourier space measurements using the same simulations (Chan, Sheth, Scoccimarro, 2012)

- measurements better described by non-local model
- convergence between 30-40 Mpc/h

*Hoffmann, Gaztanaga, Scoccimarro, Crocce in prep.